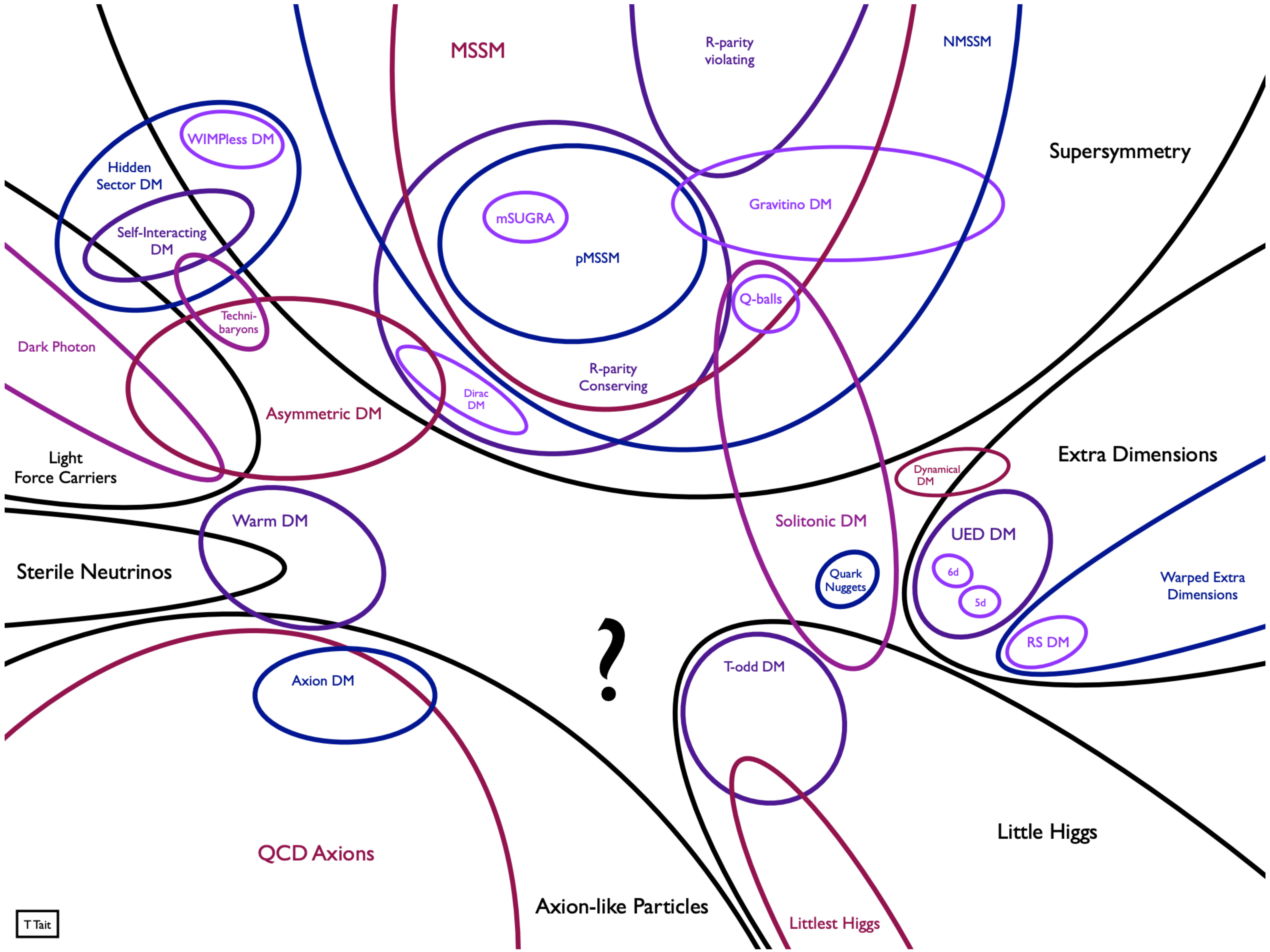
Constraints on Dark Matter Microphysics from Dwarf Galaxies

Ethan Nadler

LIneA Webinar

8/6/2020

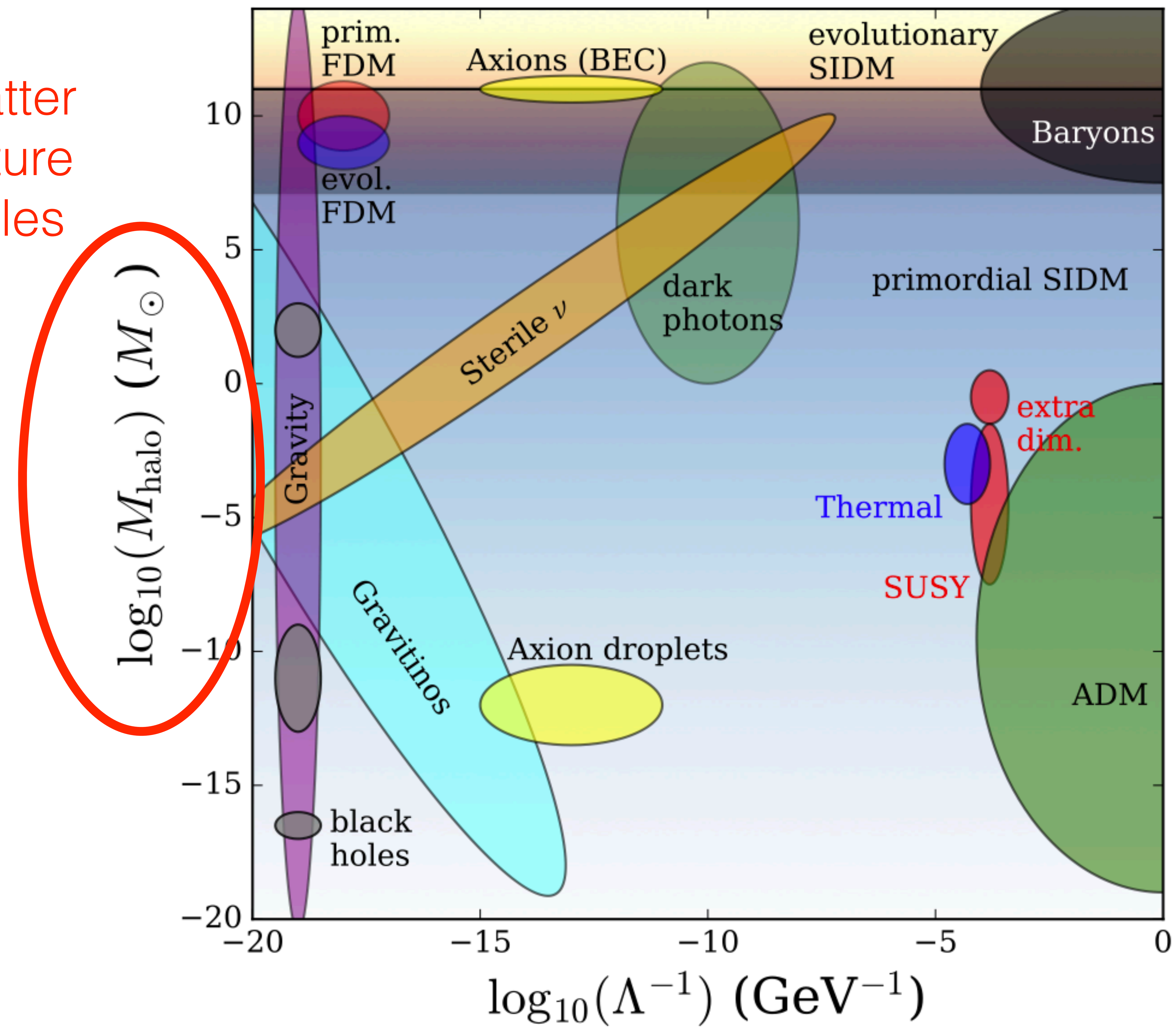
The Dark Matter Landscape



Credit: T. Tait

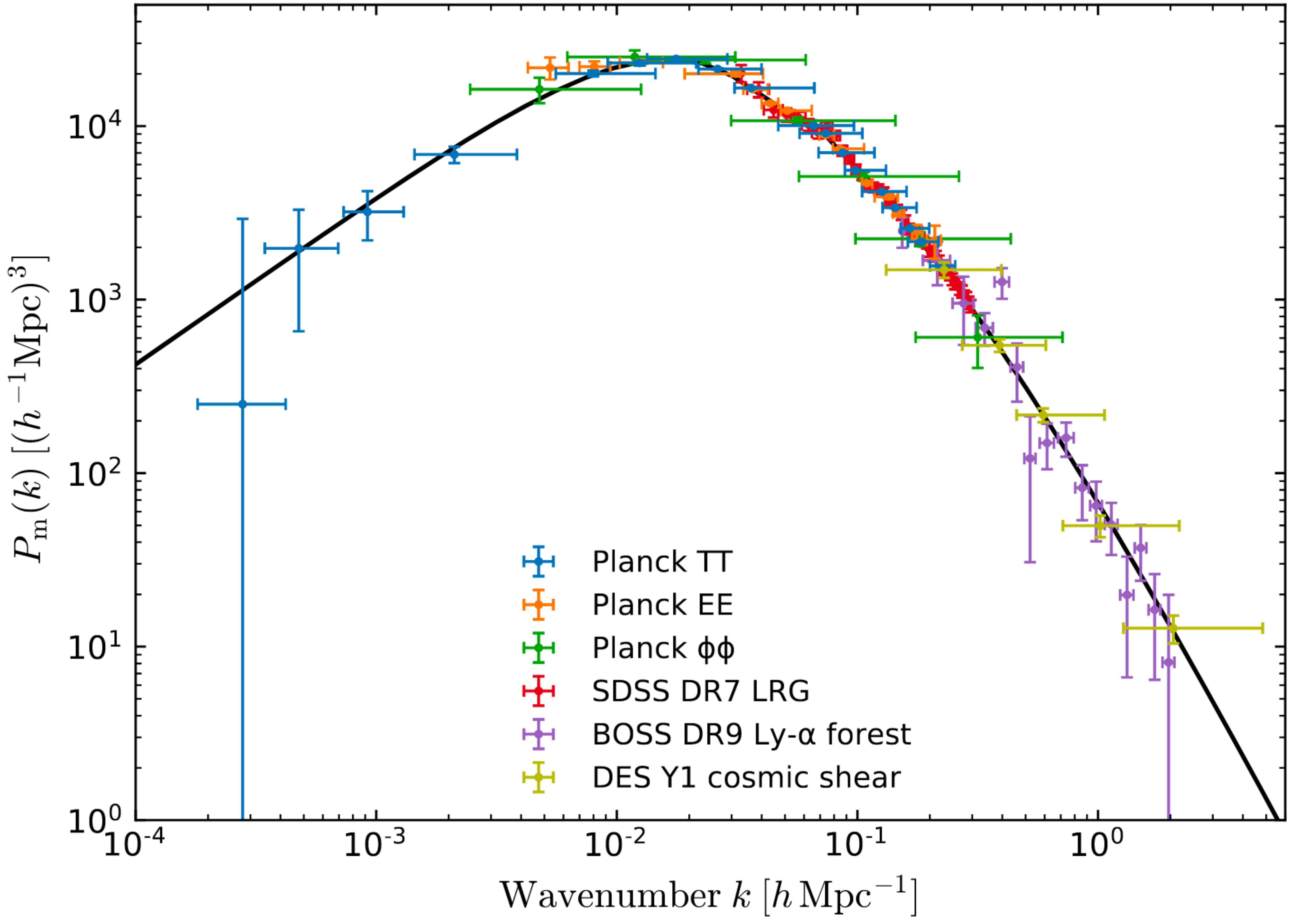
Dark Matter and Structure Formation

Microphysical dark matter properties affect structure formation on small scales

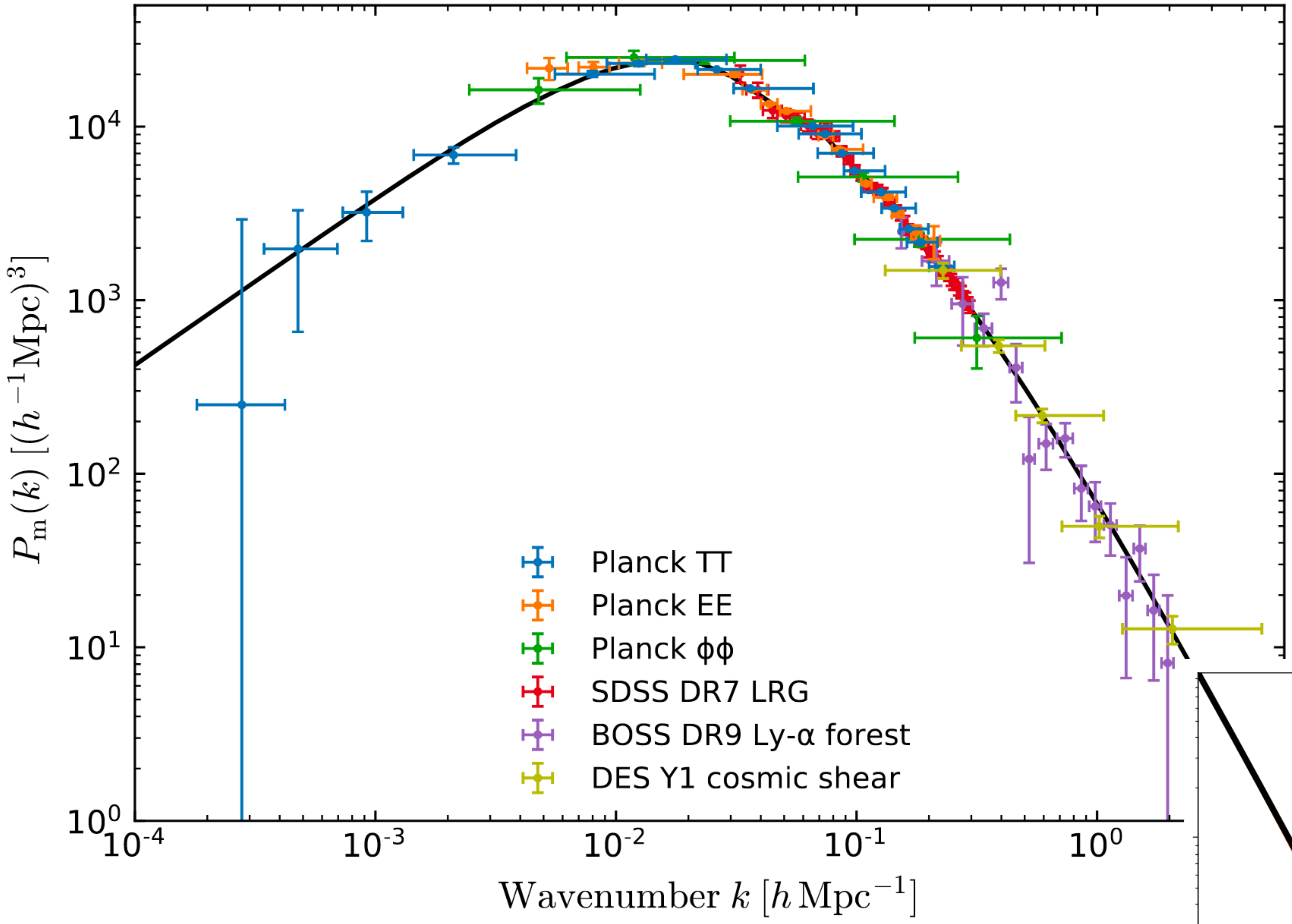


Buckley & Peter 2018

Dark Matter and Structure Formation

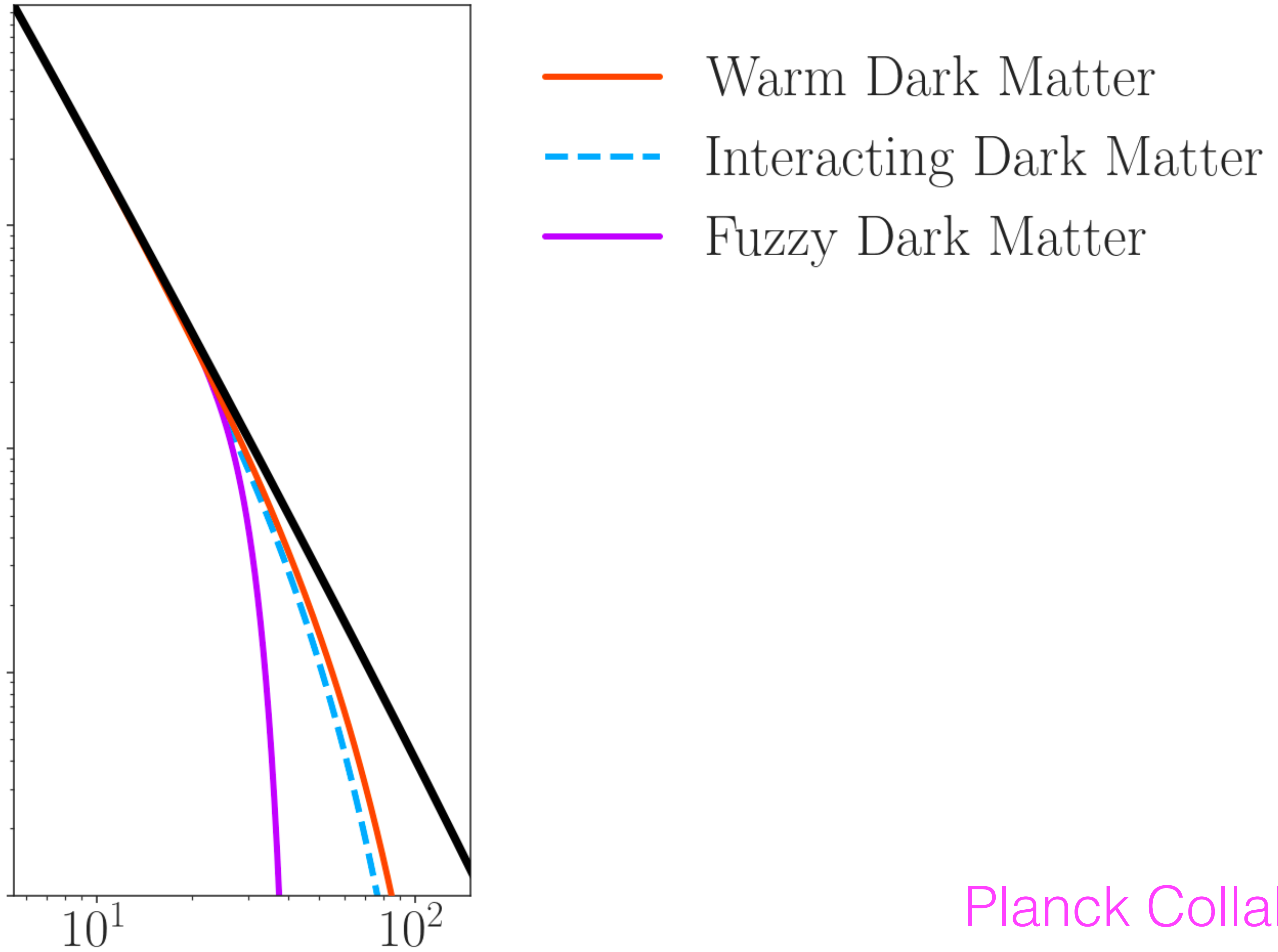


Planck Collaboration 2019



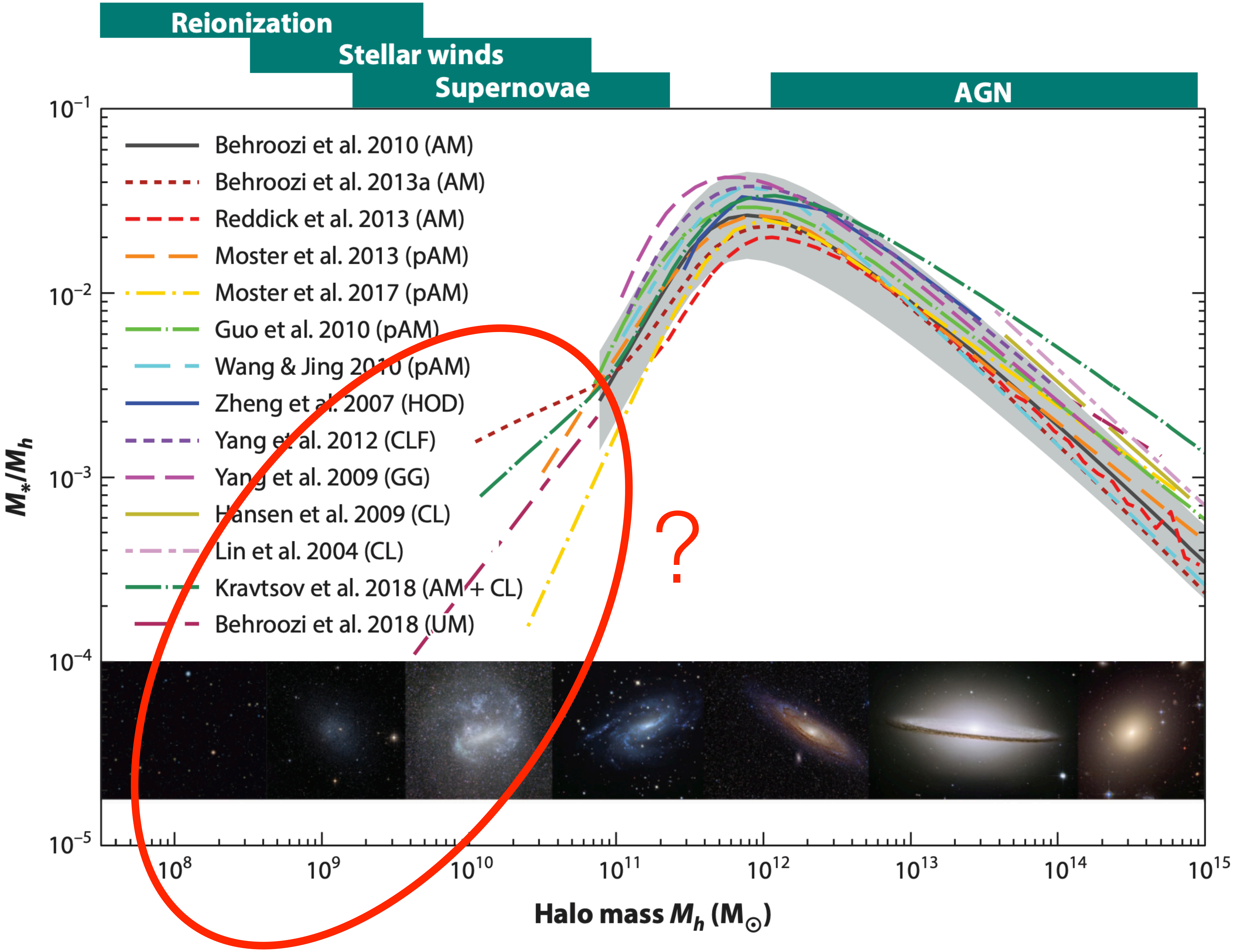
- Small scales contain information about a variety of dark matter physics: we are compelled to search there!

- **What is the luminous content of the smallest halos?**



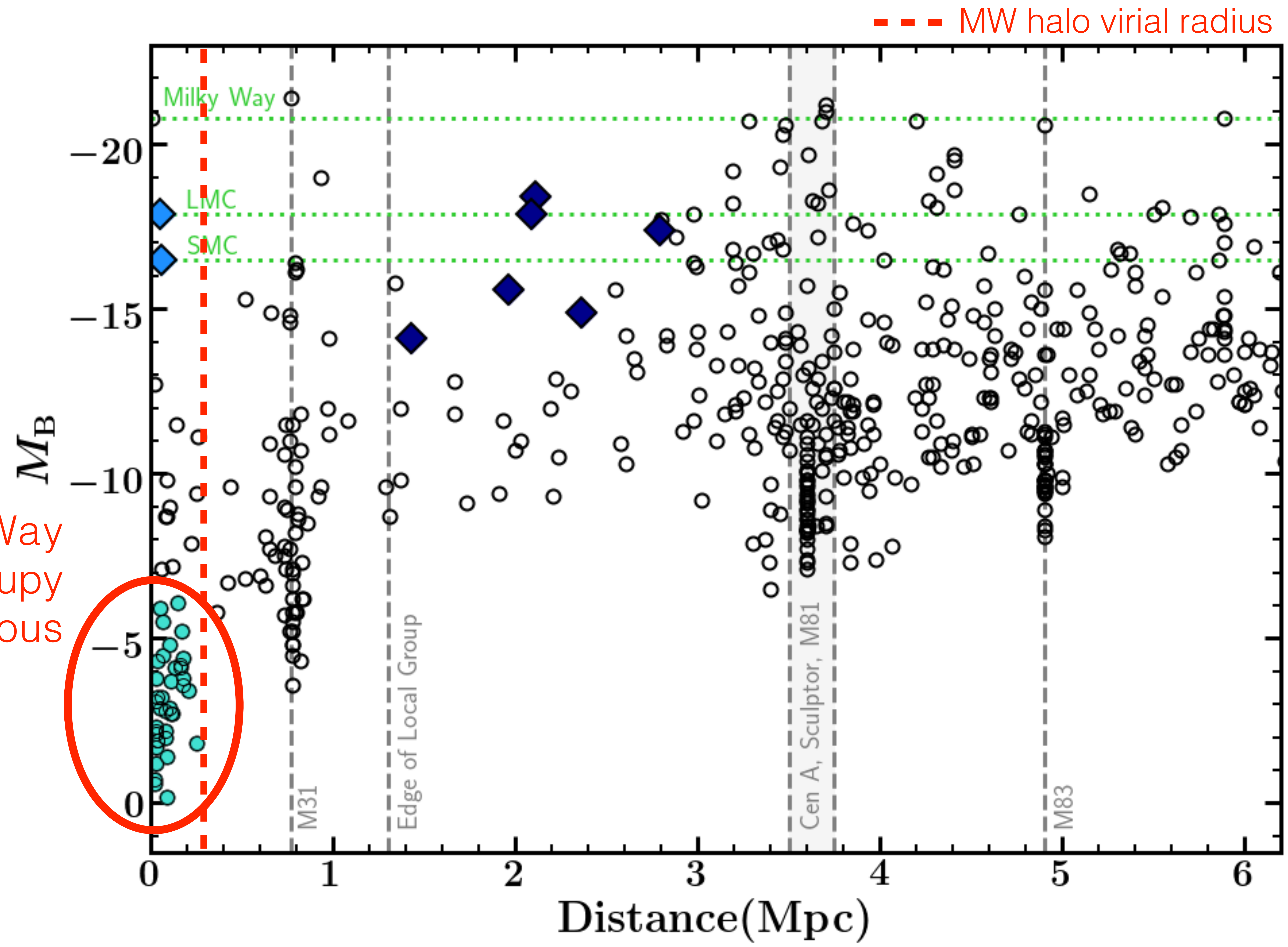
Planck Collaboration 2019

The Galaxy–Halo Connection



Wechsler & Tinker 2018

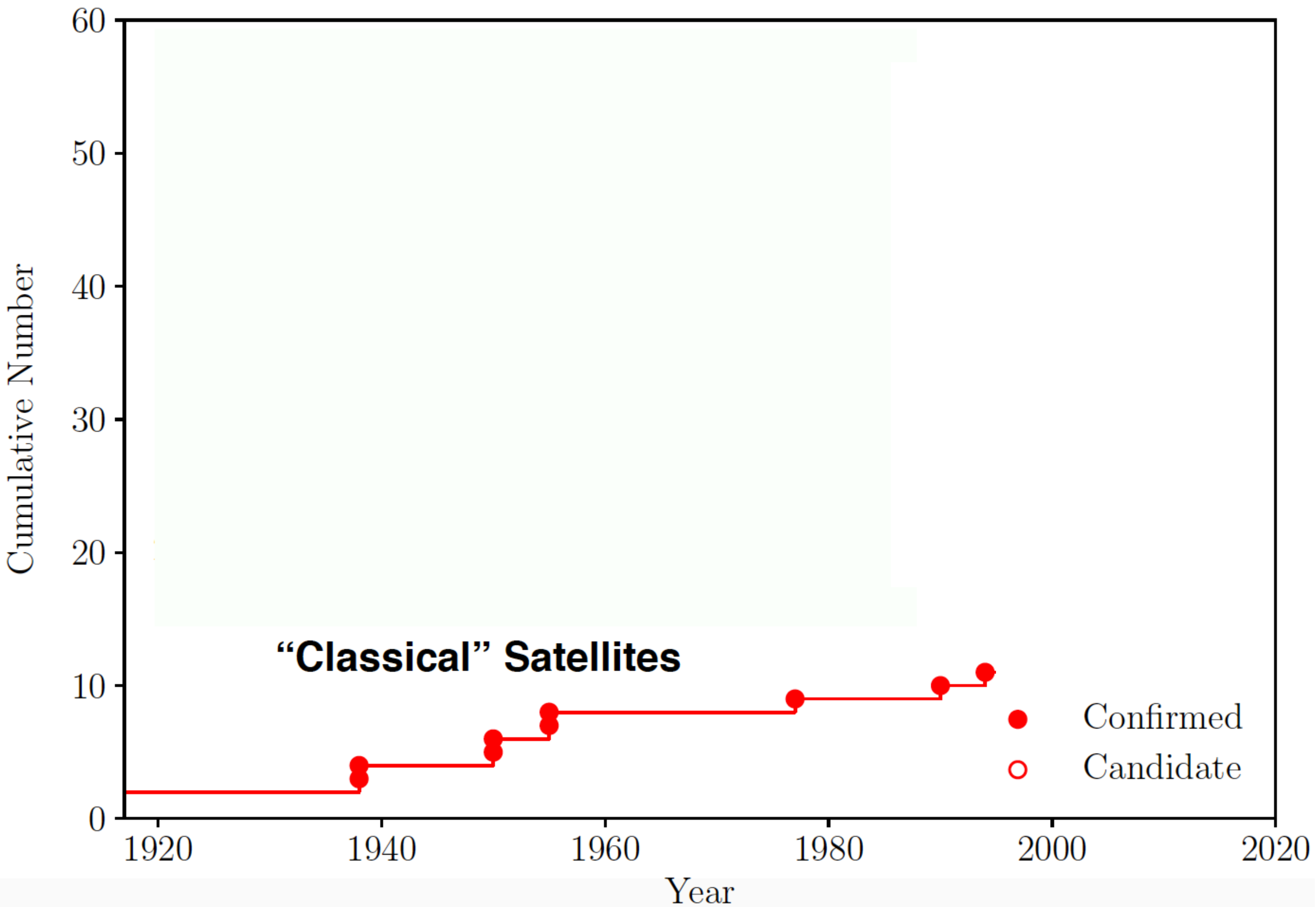
Our Census of the Faintest Galaxies



Ultra-faint Milky Way satellite galaxies occupy the smallest luminous dark matter halos

Credit: J. Carlin

The Milky Way Satellite Population



Credit: K. Bechtol

The Milky Way Satellite Population

Drlica-Wagner & Bechtol et al. 2020

Missing Satellites?

Dwarf galaxy problem

From Wikipedia, the free encyclopedia

“although there seem to be enough observed normal-sized galaxies to match the simulated size distribution, the number of dwarf galaxies is orders of magnitude lower than expected from simulation.”

There is No Missing Satellites Problem

Stacy Y. Kim, Annika H. G. Peter, Jonathan R. Hargis

“We show that there is a match between the observed satellite counts corrected by the detection efficiency of the Sloan Digital Sky Survey … and the number of luminous satellites predicted by CDM, assuming an empirical relation between stellar mass and halo mass. The “missing satellites problem”, cast in terms of number counts, is thus solved.”

Quanta Magazine

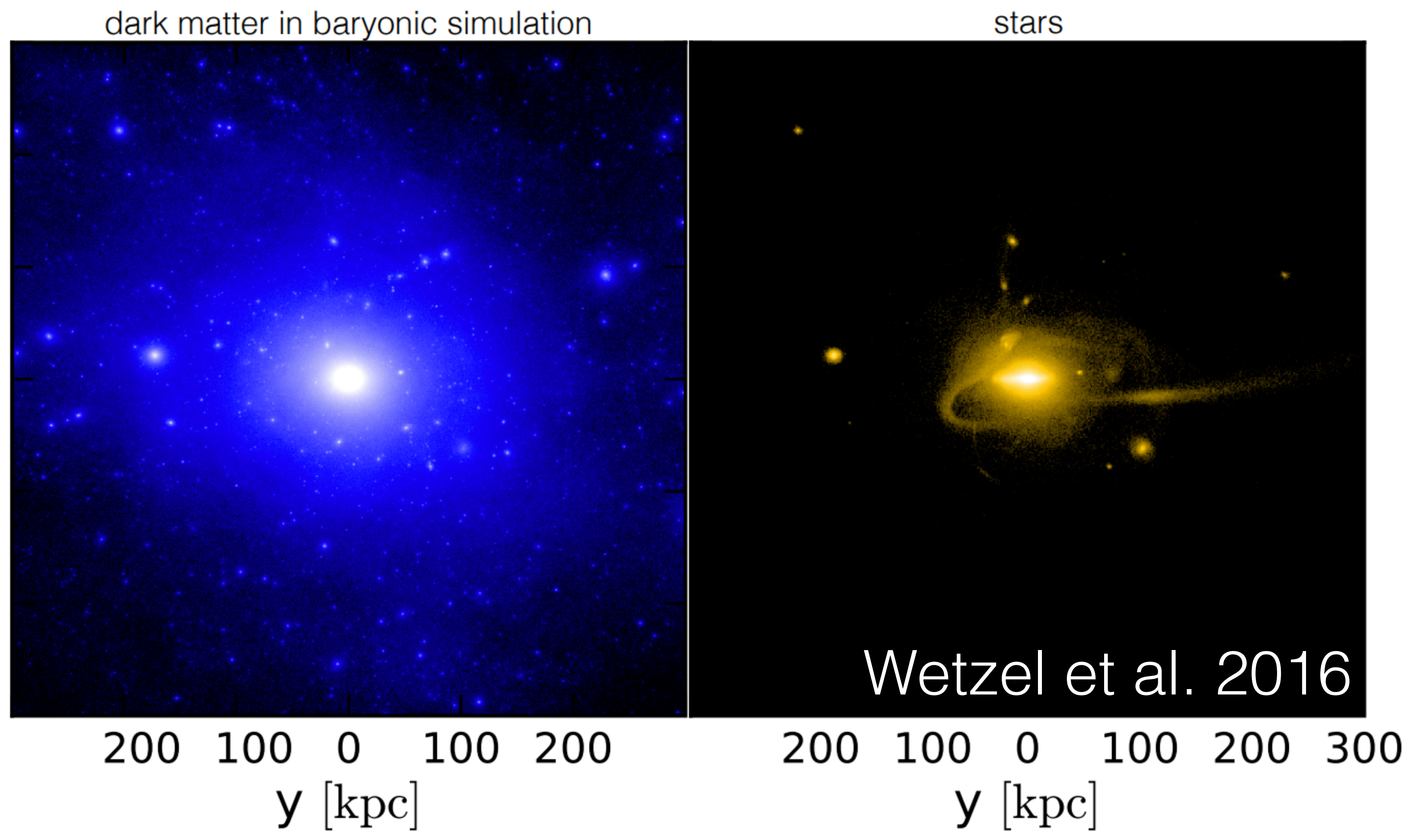
The Problem of the Missing Satellite Galaxies Gives Way — Now There's Too Many

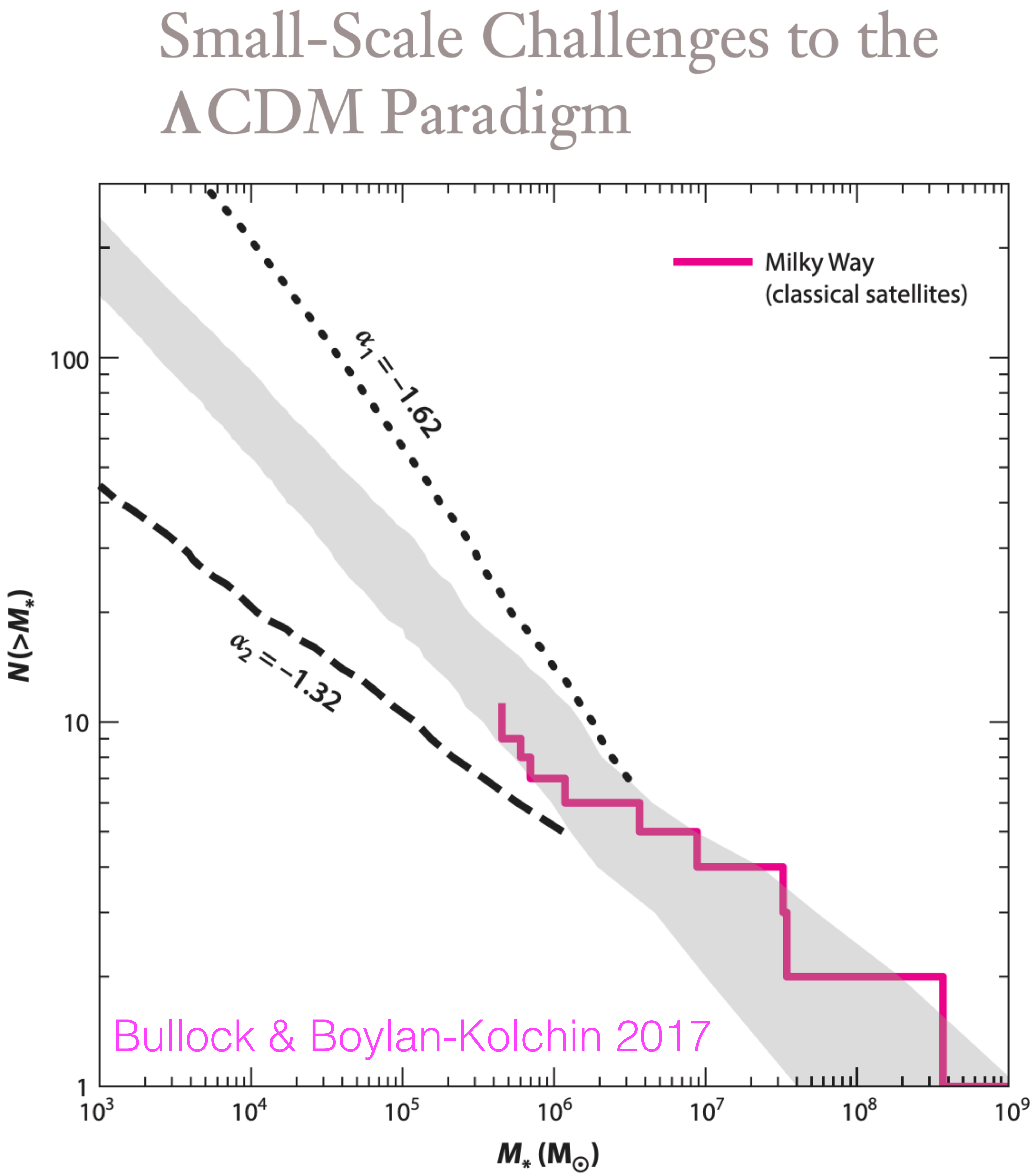
“Astronomers couldn’t find enough satellite galaxies orbiting the Milky Way. Now they have the opposite problem, suggesting that our understanding of how galaxies get built is incomplete.”

No Missing Satellites?

- No missing satellites after correcting for observational completeness and extrapolating a standard stellar mass–halo mass relation

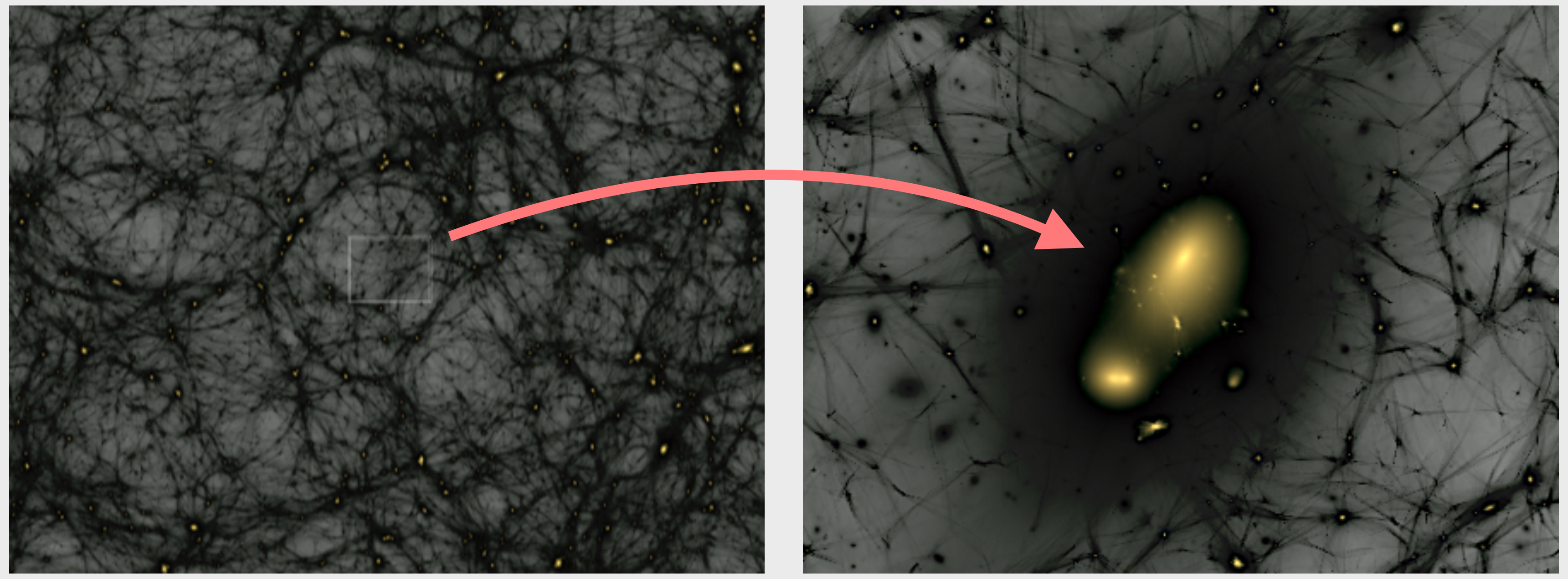
- Consistency has only been demonstrated at fixed modeling assumptions, and only for the brightest half of the observed satellites



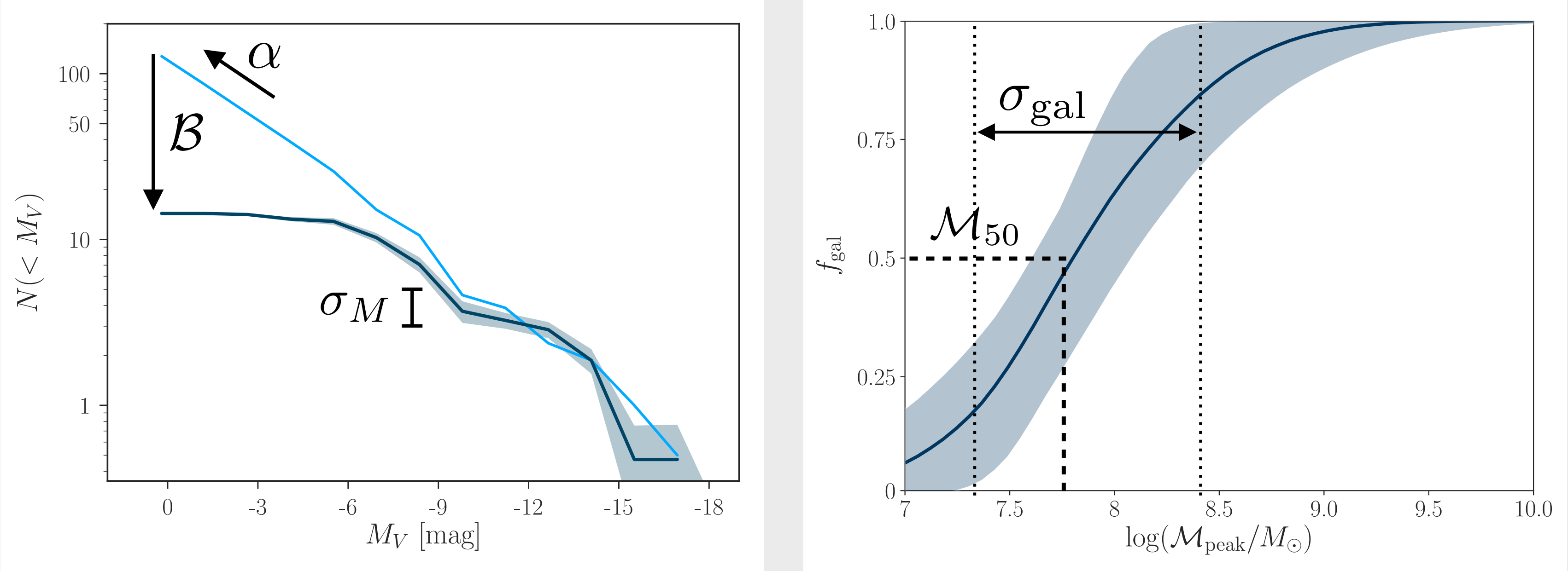


Markov Chain Monte Carlo

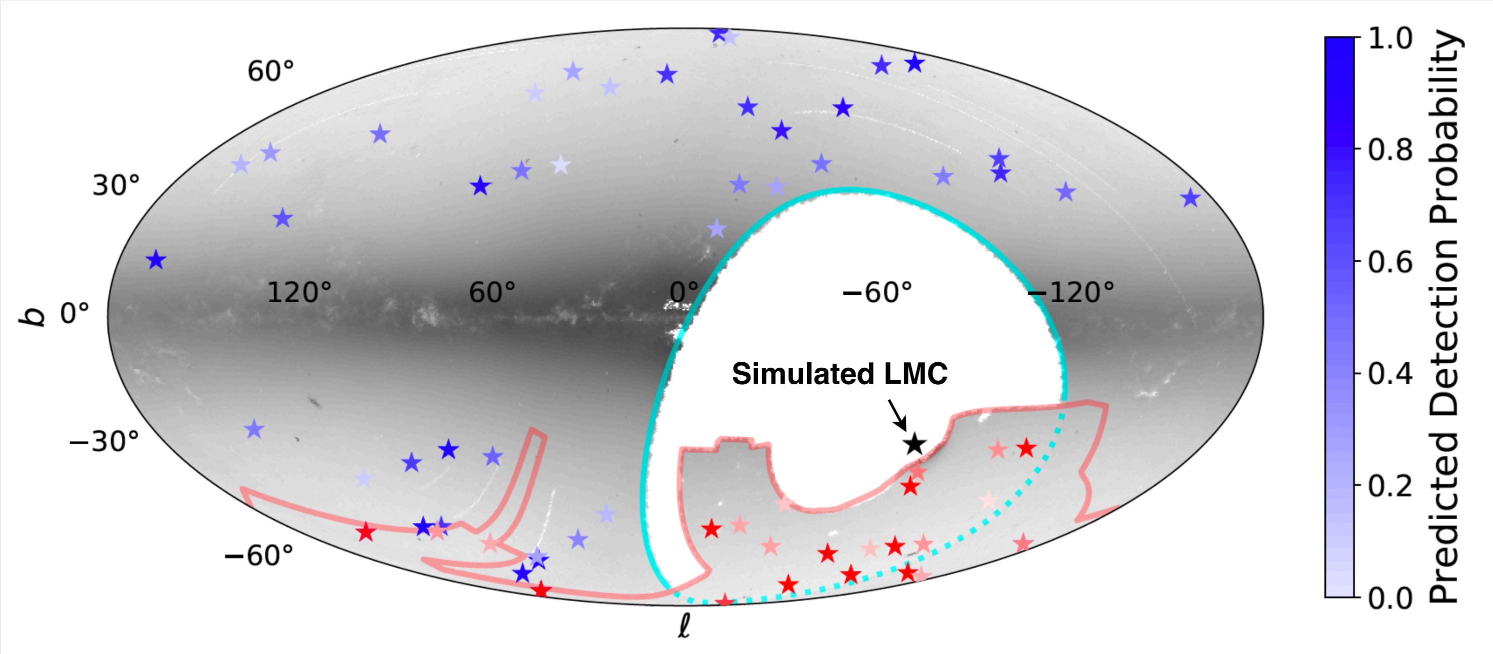
1. **Resimulate Milky Way-like halos from large cosmological volume.**



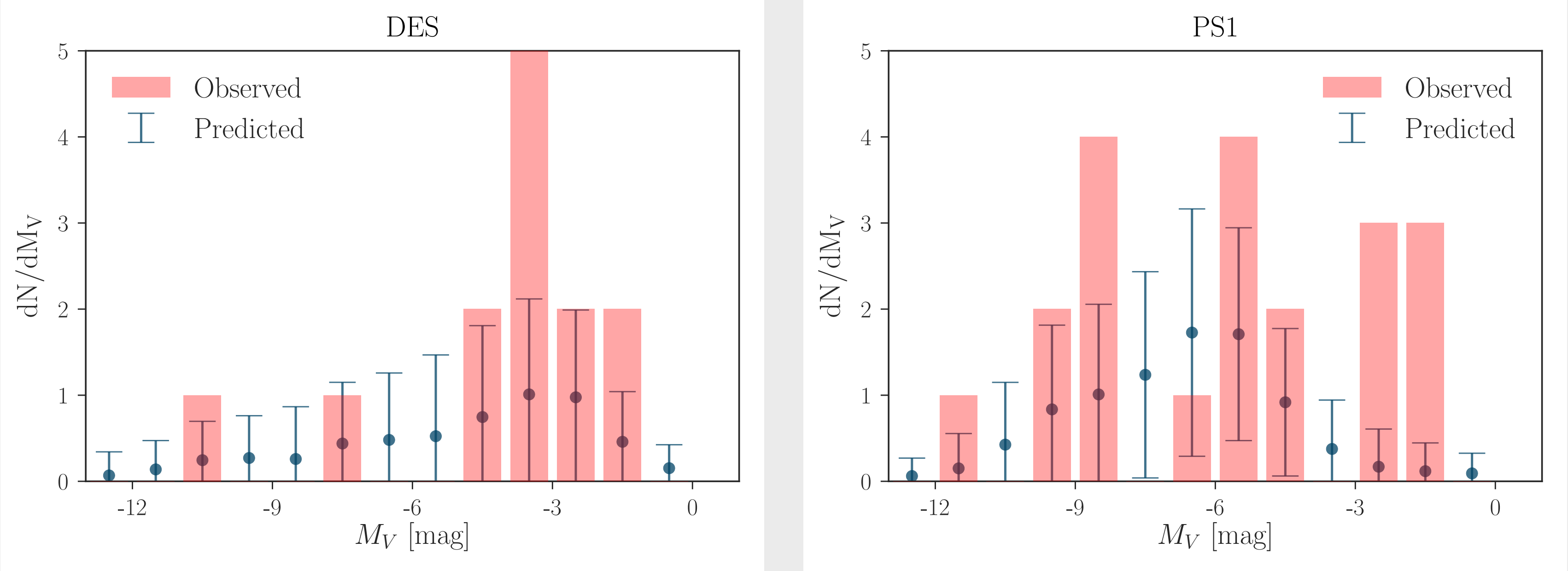
2. **Paint satellite galaxies onto subhalos using galaxy—halo model.**



3. **Apply observational selection functions based on imaging data.**



4. **Calculate likelihood of observed satellites given galaxy—halo connection parameters.**

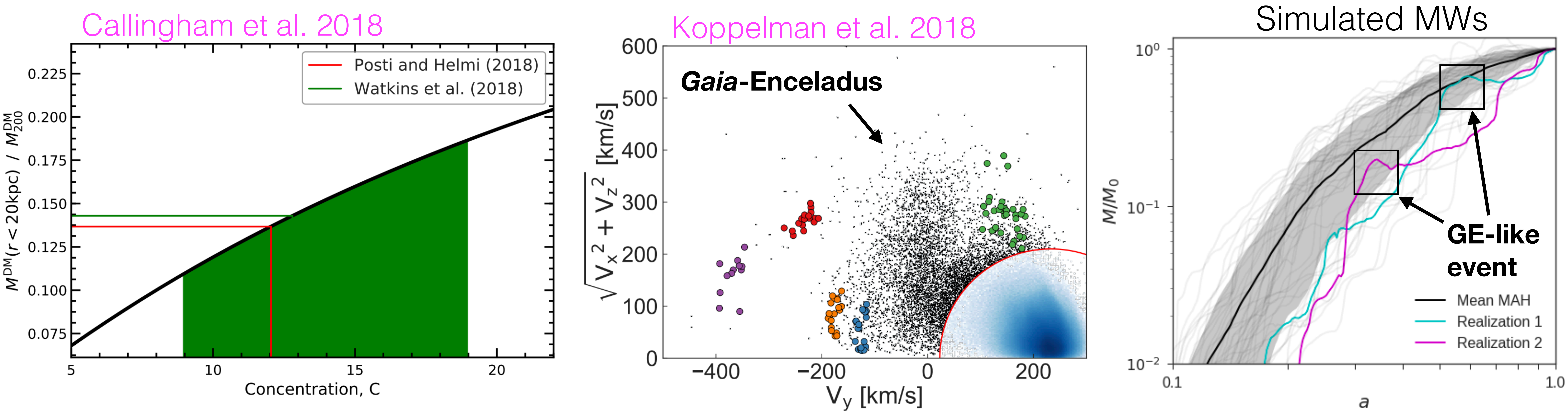


EN & Wechsler et al. 2020

Simulating Milky Way Analogs

- *Gaia* is revolutionizing our understanding of Milky Way halo properties: mass, concentration, assembly history

- We analyze halos that 1) resemble the MW in these properties, 2) experience a *Gaia*-Enceladus merger, and 3) have a realistic Large Magellanic Cloud analog



Gaia-Enceladus
LMC
MW
Credit: Ralf Kaehler

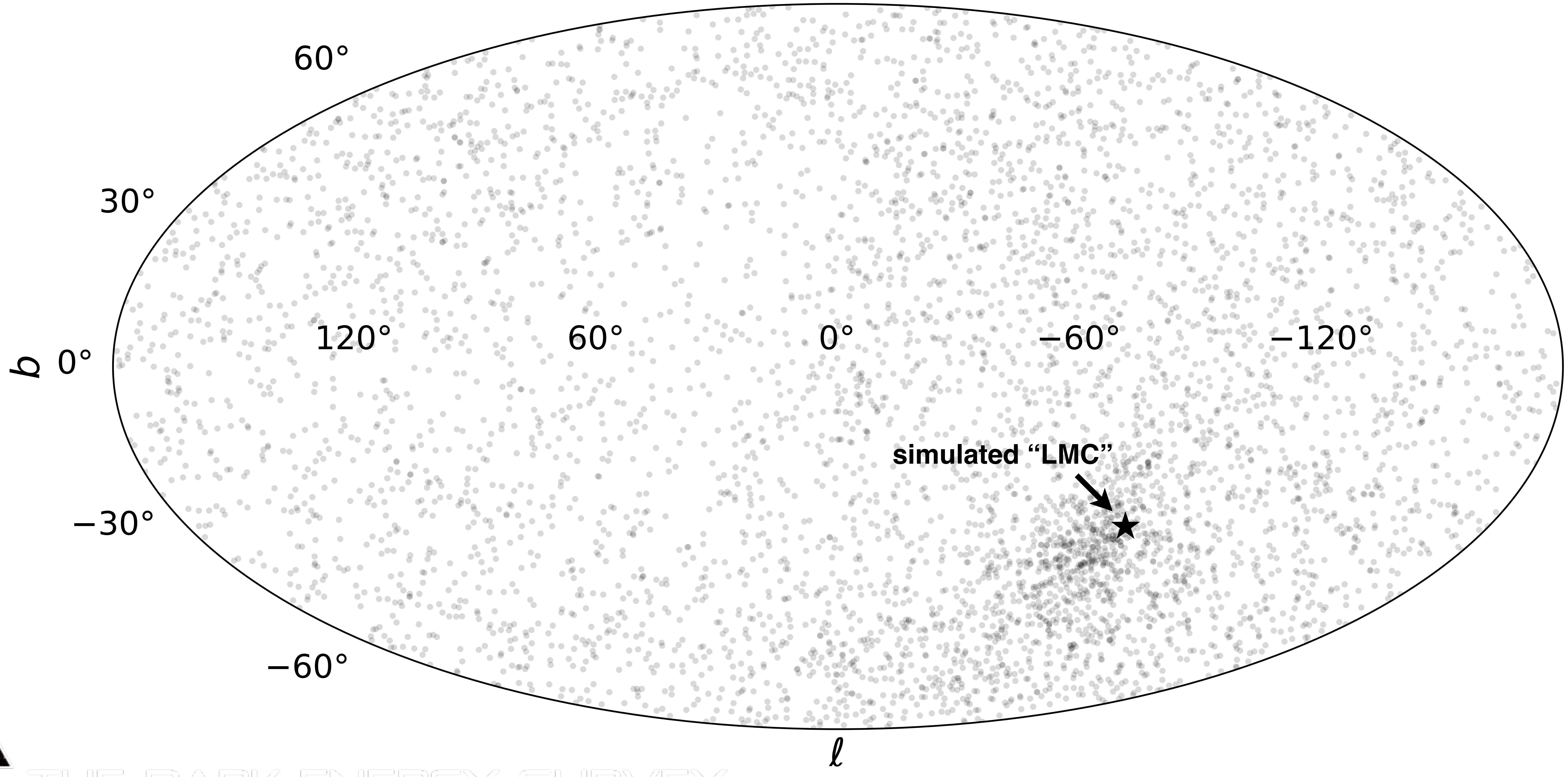
Galaxy–Halo Connection Model

Empirical modeling allows us to marginalize over theoretical uncertainties

Physical Ingredient	Assumptions	Parameterization	Free Parameter?
Satellite Luminosities	Abundance match to GAMA survey	Non-parametric	*No*
	Extrapolate luminosity function	Faint-end slope α	**Yes** (α is free)
	Lognormal $(M_V \vert V_{\rm peak})$ distribution	Constant scatter σ_M	**Yes** (σ_M is free)
	Smooth galaxy formation efficiency	$f_{\rm gal} \equiv \frac{1}{2}\left[1+\left(\frac{\mathcal{M}_{\rm peak}-\mathcal{M}_{50}}{\sqrt{2}\sigma_{\rm gal}}\right)\right]$	**Yes** ($\mathcal{M}_{50}, \sigma_{\rm gal}$ are free)
Satellite Sizes	Kravtsov (2013) galaxy size model	$r_{1/2} \equiv \mathcal{A}\,(R_{\rm vir}/R_0)^n$	**Yes** ($\mathcal{A}, n$ are free)
	Lognormal $(r'_{1/2} \vert R_{\rm vir})$ distribution	Constant scatter σ_R	**Yes** (σ_R is free)
	Size reduction set by stripping	$r'_{1/2} \equiv r_{1/2}\,(V_{\rm max}/V_{\rm acc})^{\beta}$	*No* ($\beta = 0$)
Baryonic Effects	Nadler et al. (2018) disruption model	$p_{\rm disrupt} \rightarrow p_{\rm disrupt}^{1/\mathcal{B}}$	**Yes** ($\mathcal{B}$ is free)
Orphan Satellites	Correspond to disrupted subhalos	None	*No*
	NFW host + dynamical friction	$\ln\Lambda = -\ln(m_{\rm sub}/M_{\rm host})$	*No*
	Stripping after pericentric passages	$\dot{m}_{\rm sub} \sim -\frac{m_{\rm sub}}{\tau_{\rm dyn}}\left(\frac{m_{\rm sub}}{M_{\rm host}}\right)^{0.07}$	*No*
	$p_{\rm disrupt}$ set by time since accretion	$p_{\rm disrupt} \equiv (1-a_{\rm acc})^{\mathcal{O}}$	*No* ($\mathcal{O} = 1$)

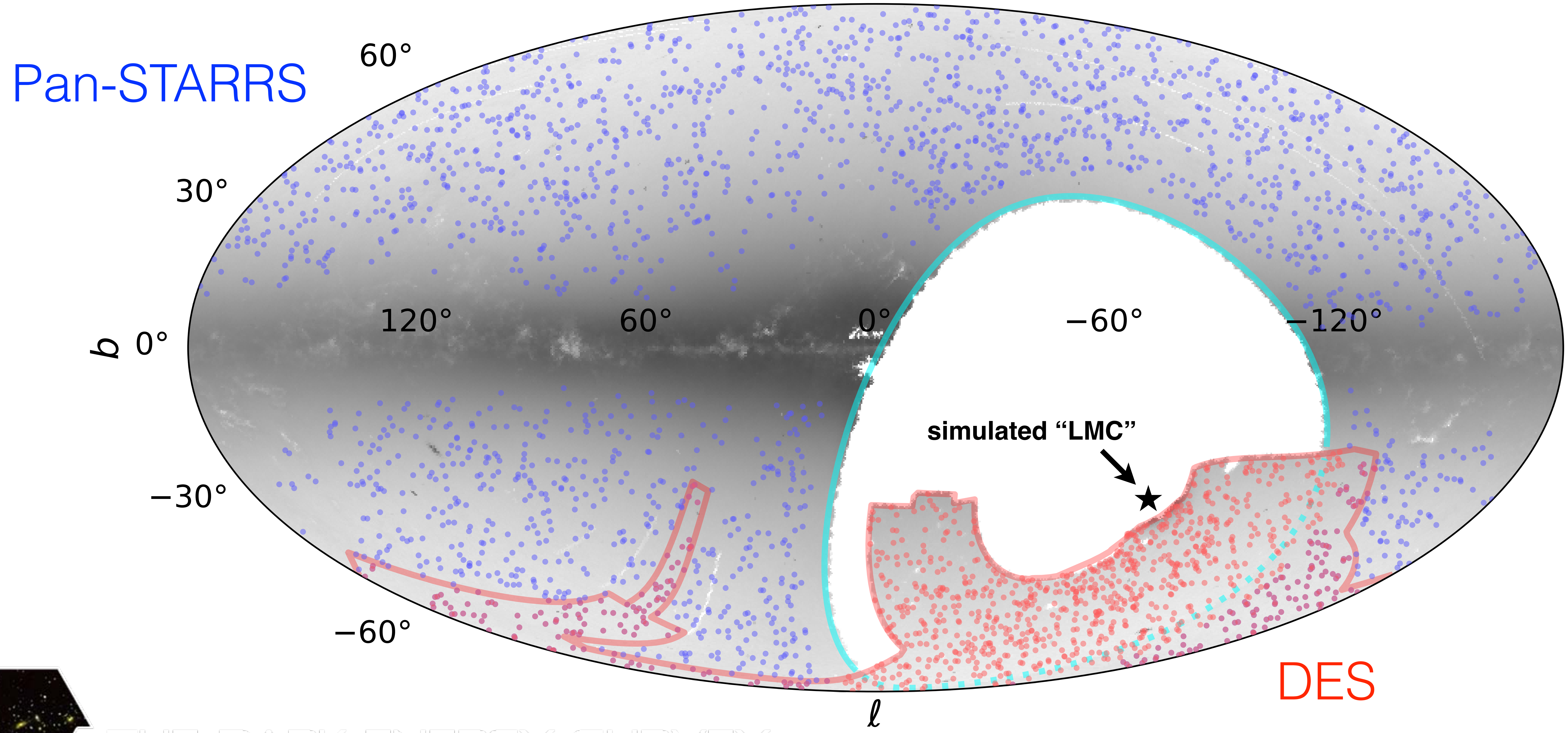
EN, Y.-Y. Mao, G. Green, R. Wechsler 2019

Mock Satellite Observations



THE DARK ENERGY SURVEY

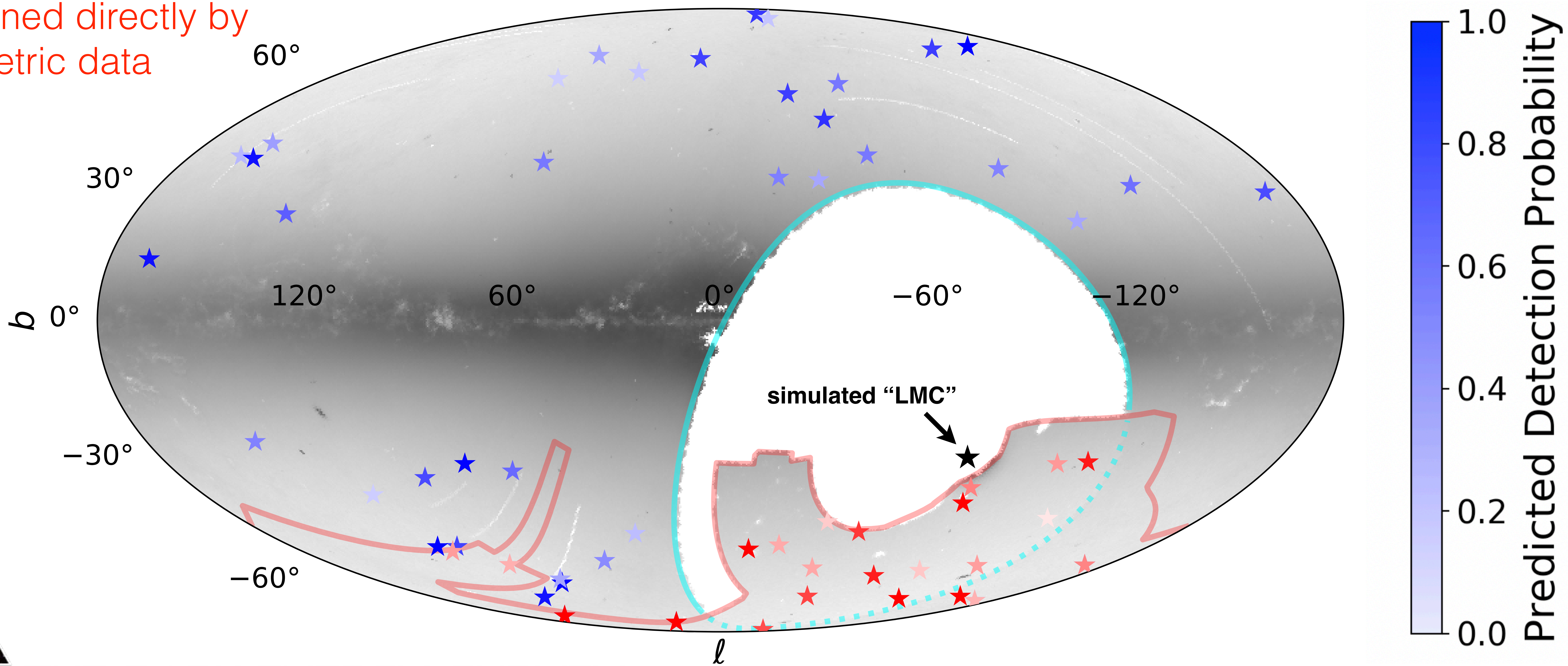
Mock Satellite Observations



THE DARK ENERGY SURVEY

Mock Satellite Observations

Detection probability is constrained directly by photometric data

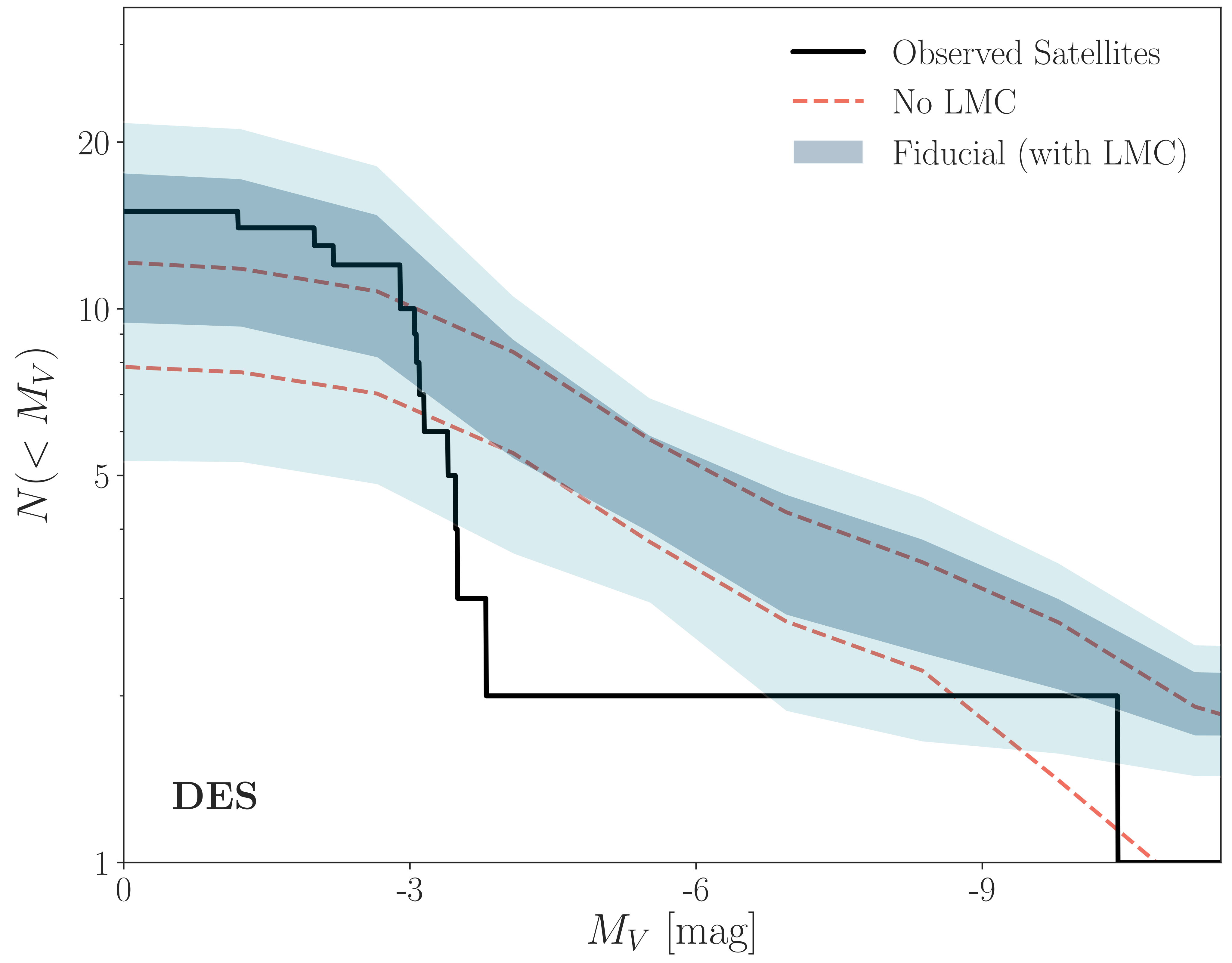


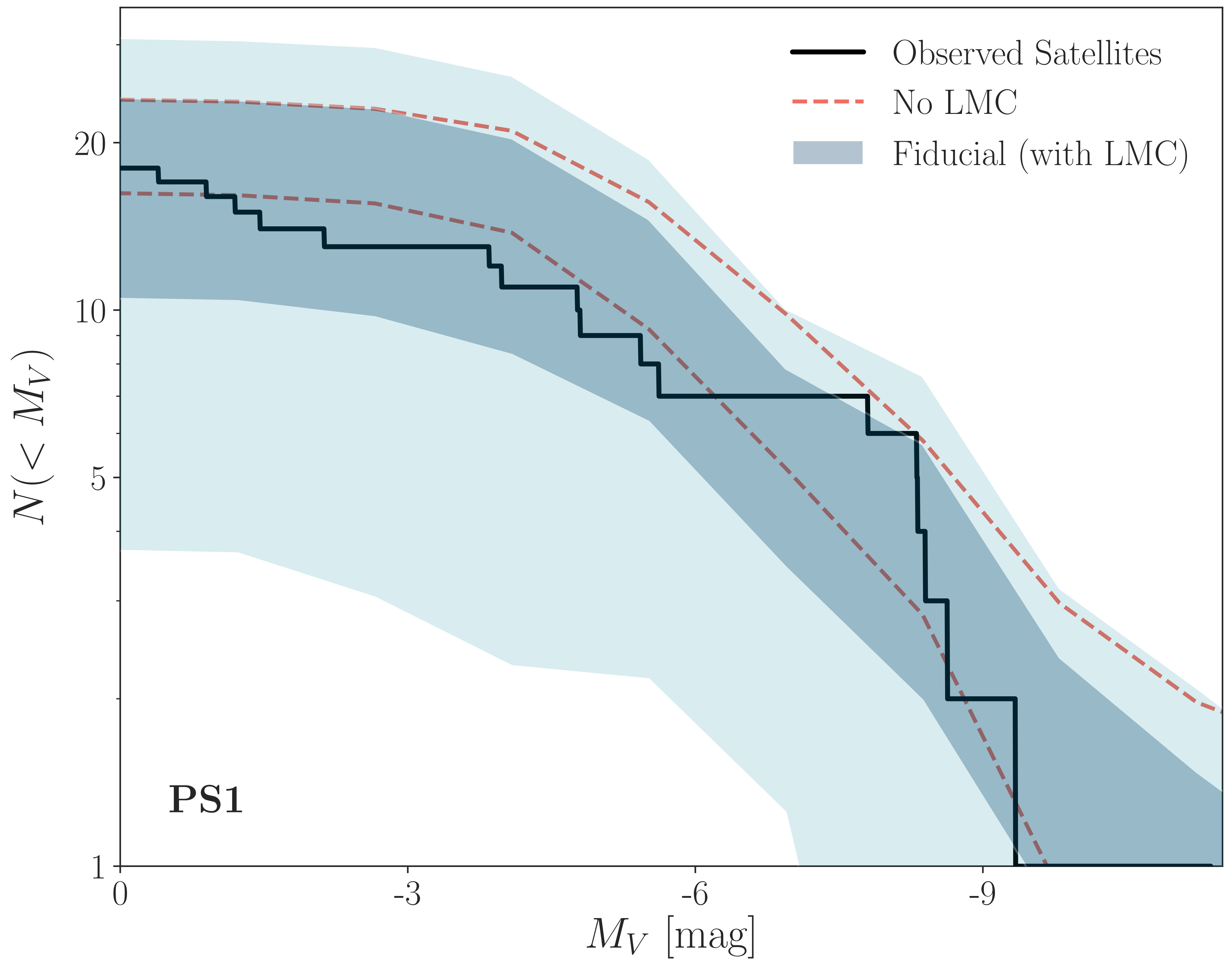
THE DARK ENERGY SURVEY

Allows for a rigorous statistical comparison to observations

Predicted Satellite Populations

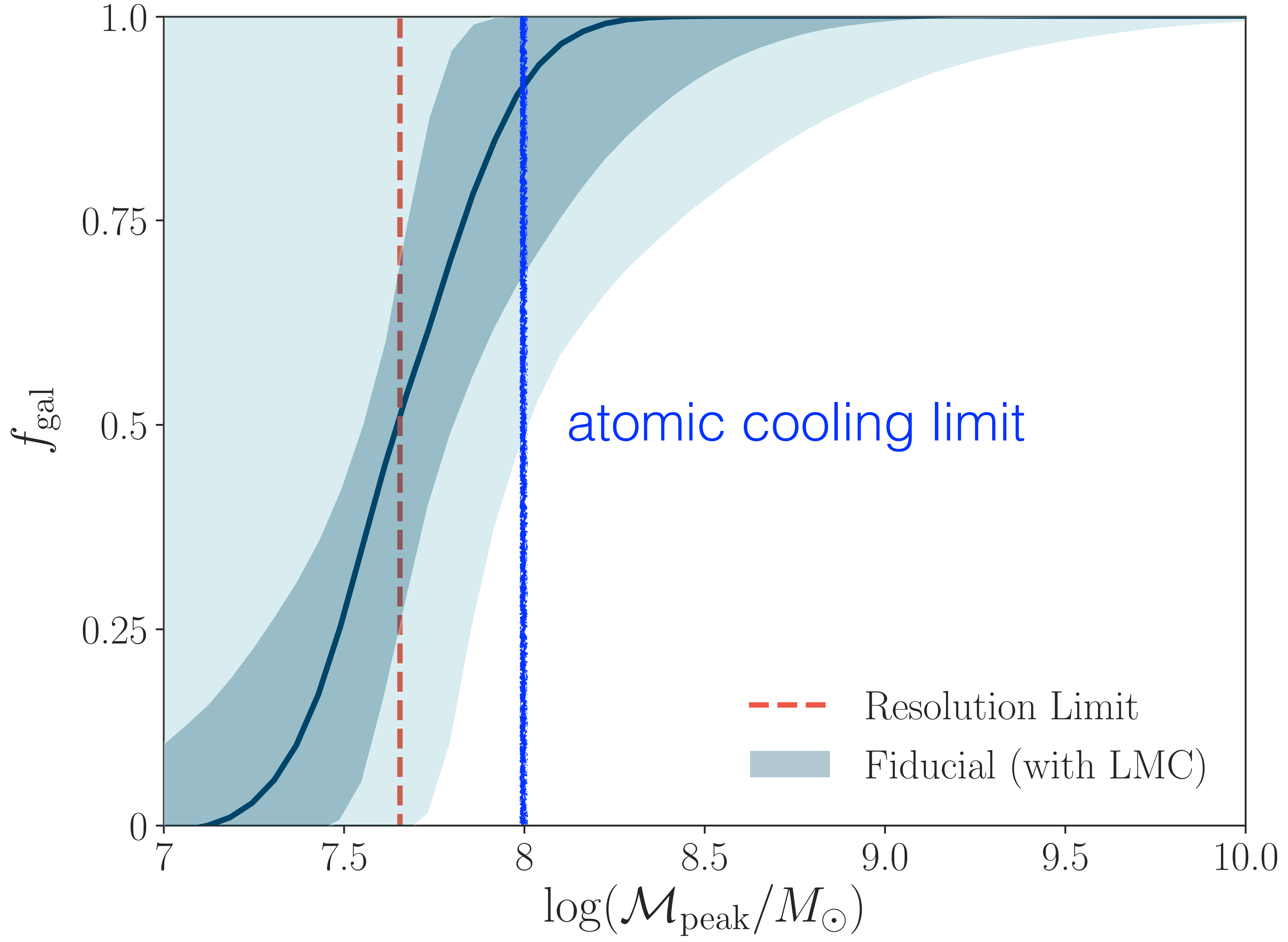
- Recent infall (within last 2 Gyr) of LMC system **required** to fit the data
- Predict 5 (DES) & 1 (PS1) LMC satellites, consistent with *Gaia* proper motions!

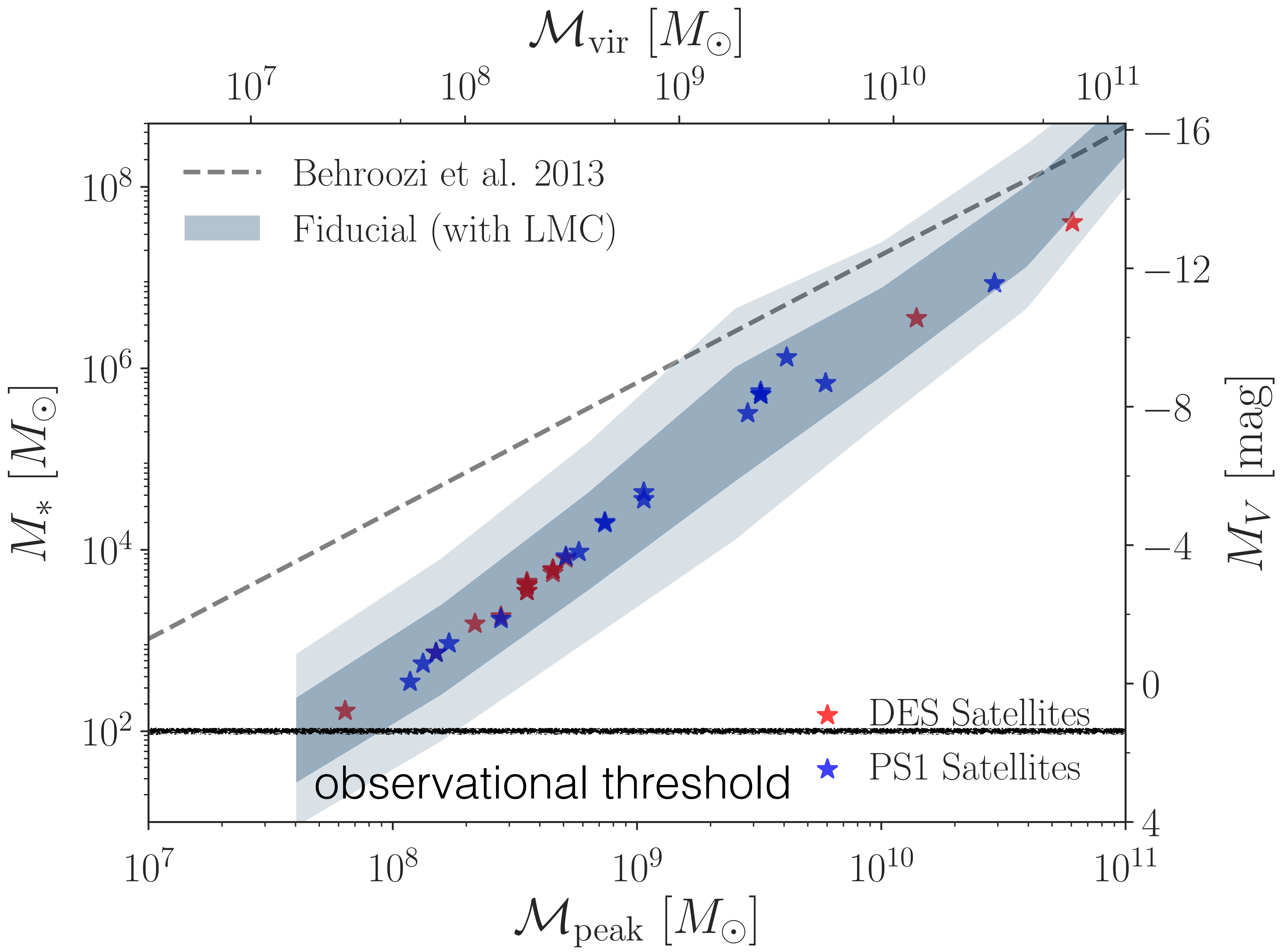




EN & Wechsler et al. 2020

The Faint-End Galaxy–Halo Connection

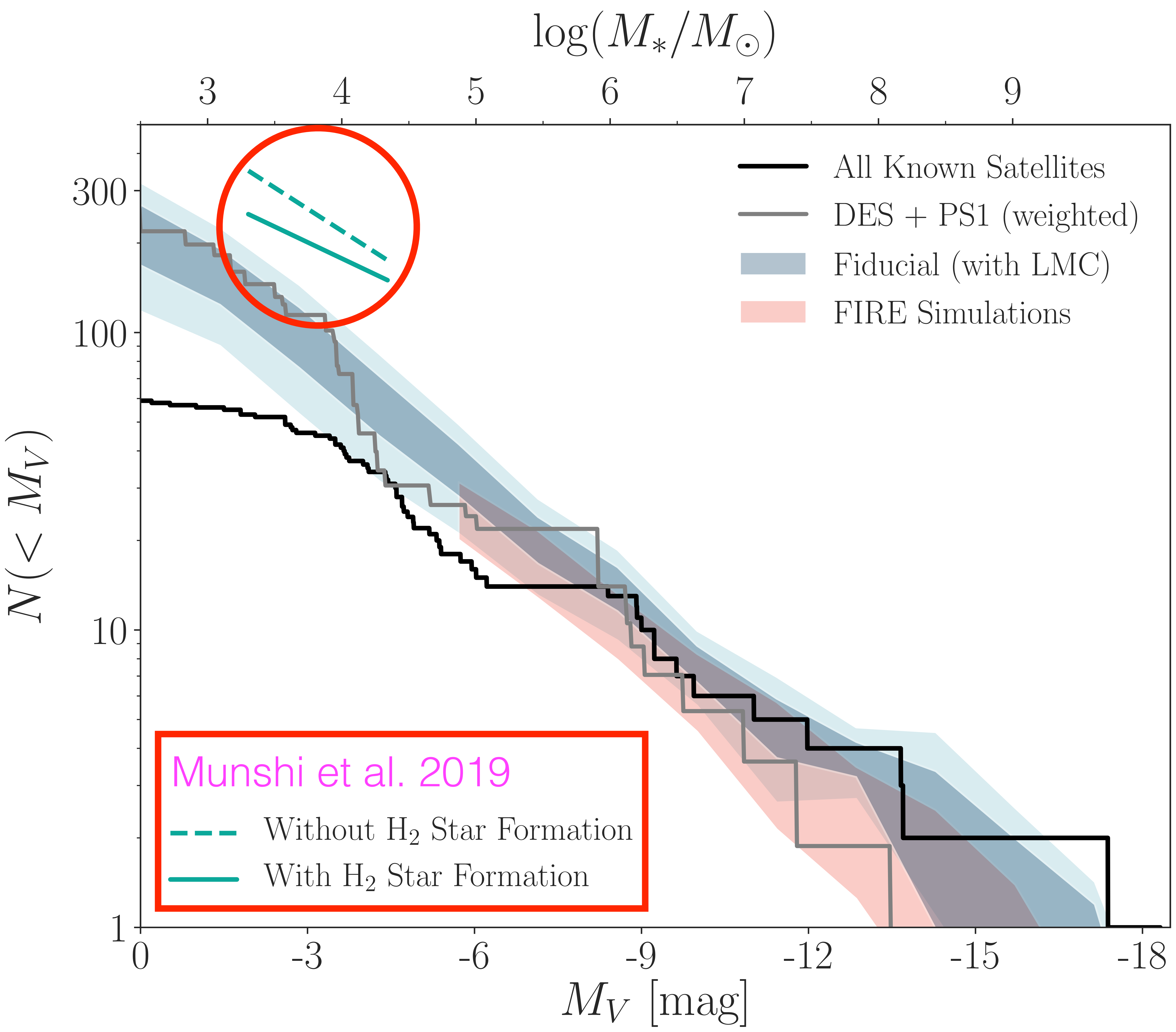




No evidence for a galaxy formation cutoff due to reionization or dark matter physics

The Total MW Satellite Population

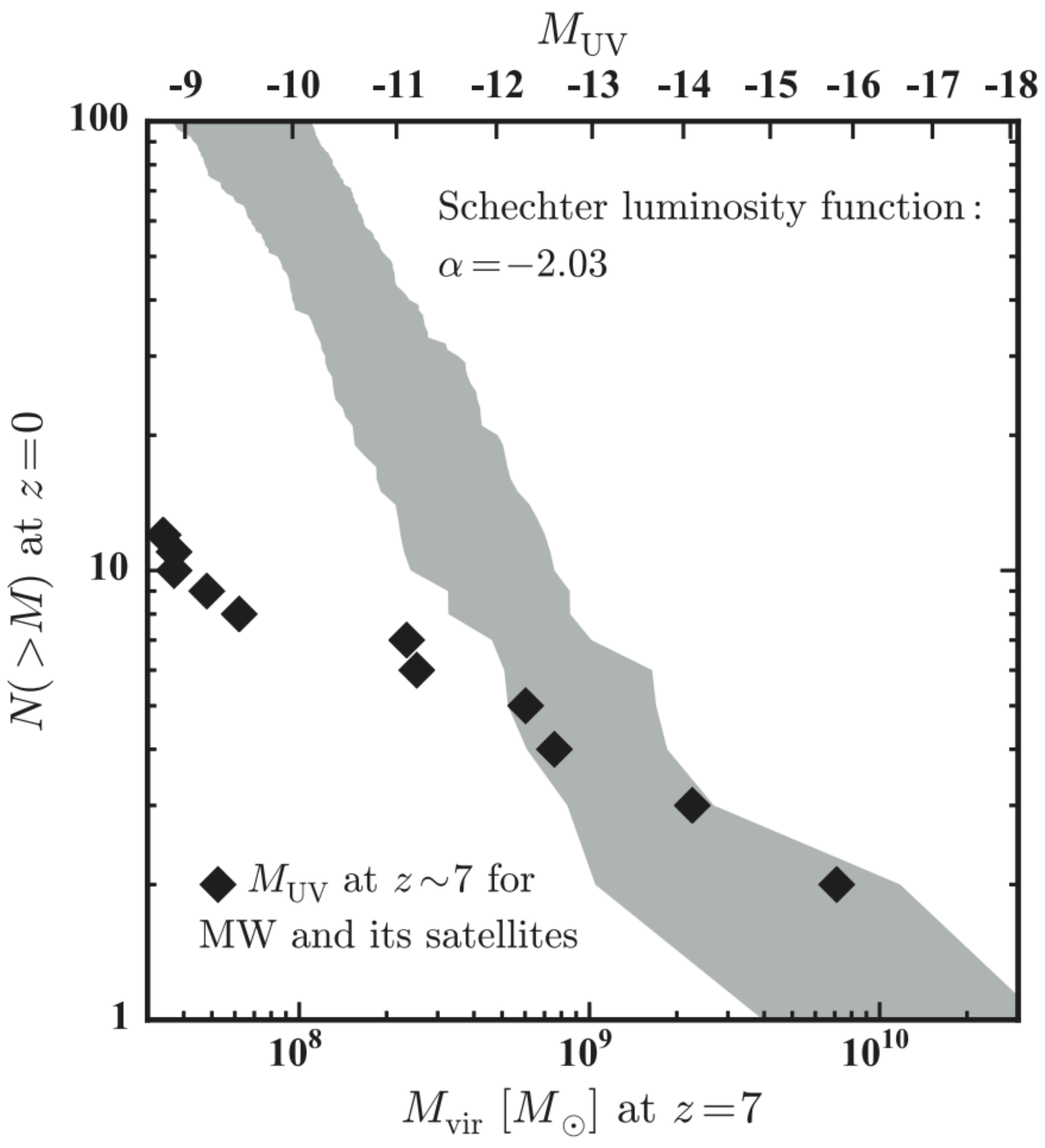
- Predict ~220 total MW satellites, 25% of which are LMC-associated
- Vera C. Rubin Observatory should discover ~100 new systems
- Spectroscopic follow-up (e.g., on GSMTs) will be key
- Constraints on faint-end slope and galaxy formation efficiency inform hydrodynamic simulations

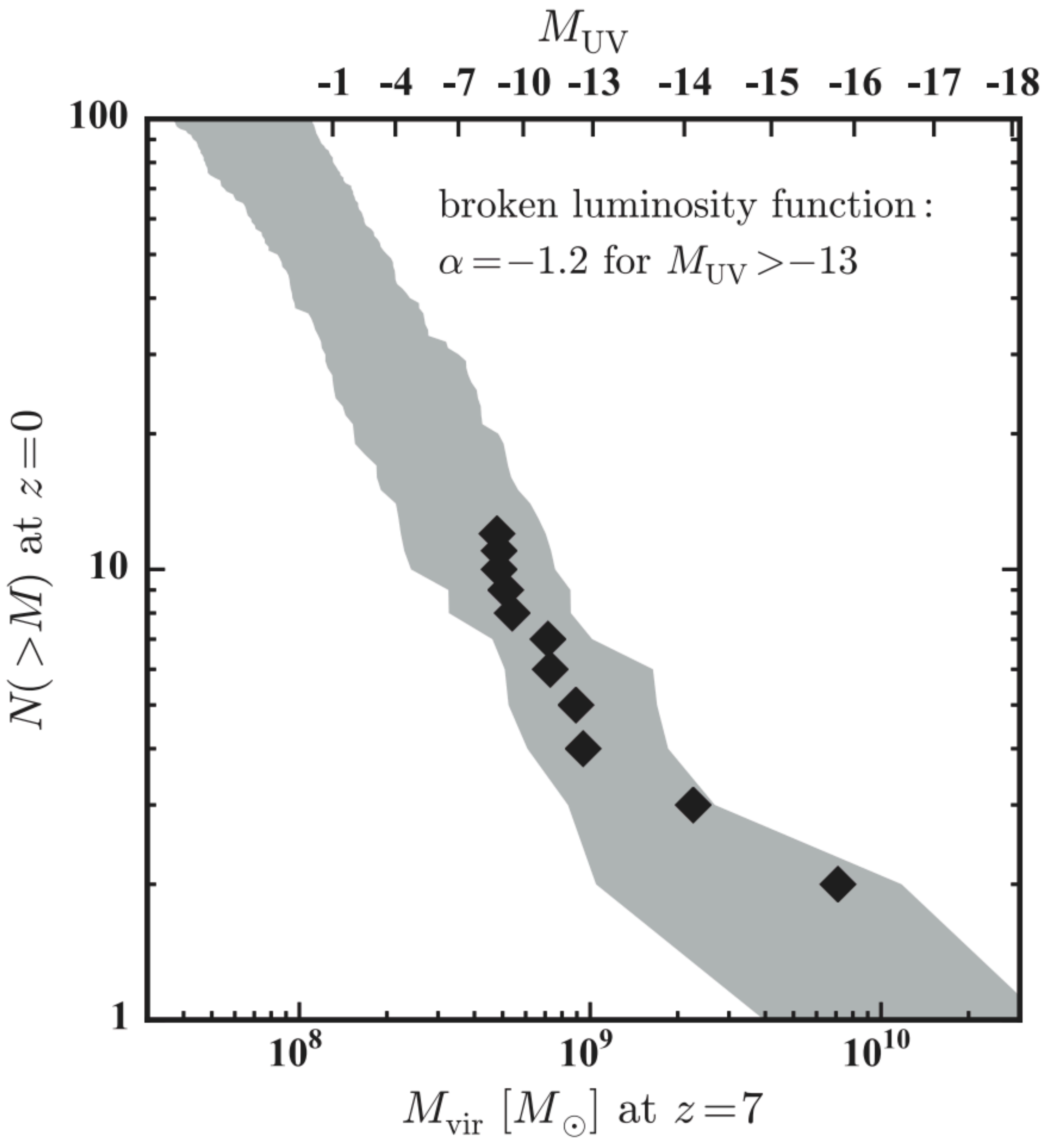


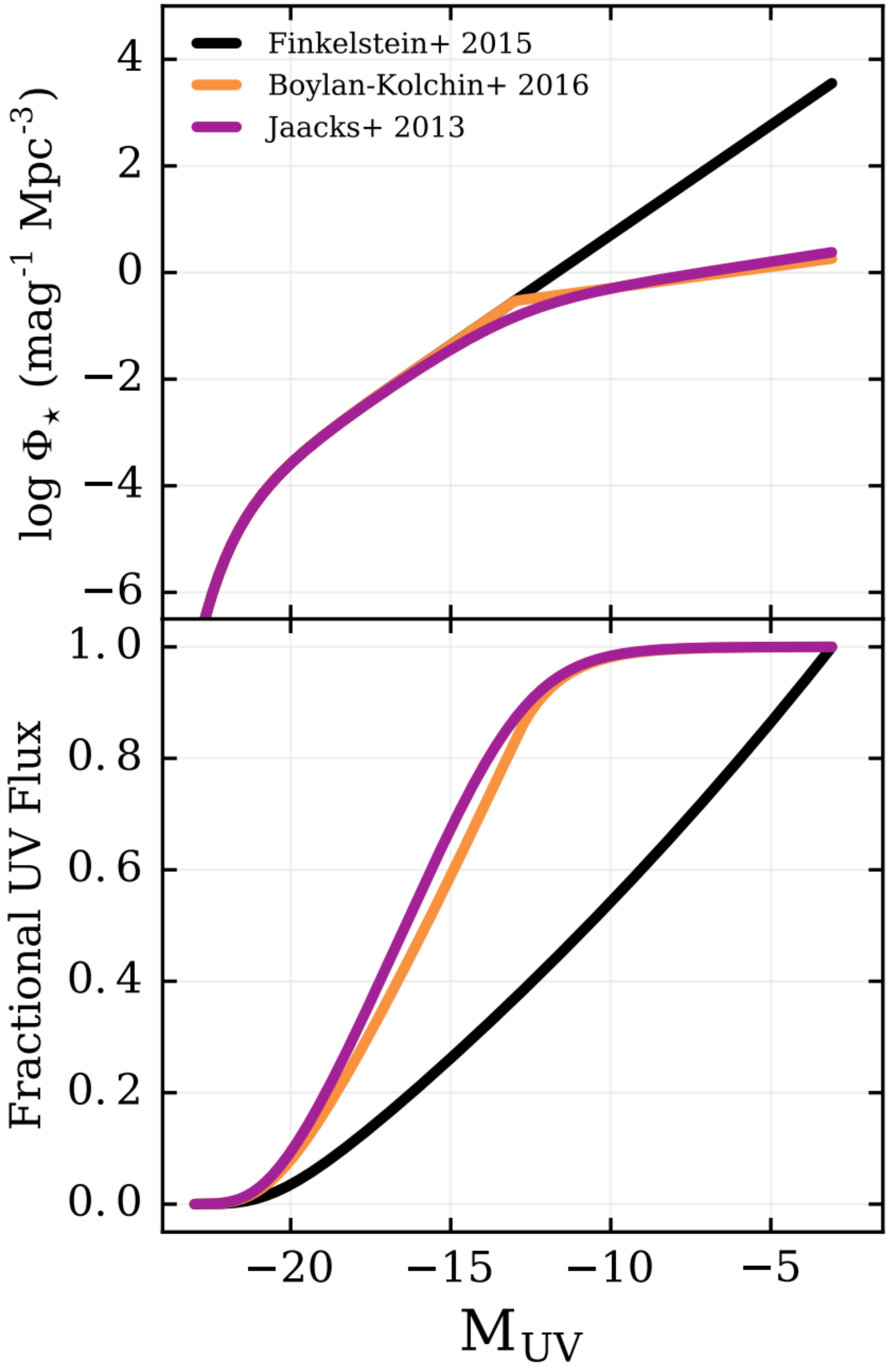
EN & Wechsler et al. 2020

Interplay with High-z Galaxy Formation

- Combining the star formation histories of Local Group dwarfs with our abundance matching results will constrain the high-redshift luminosity function





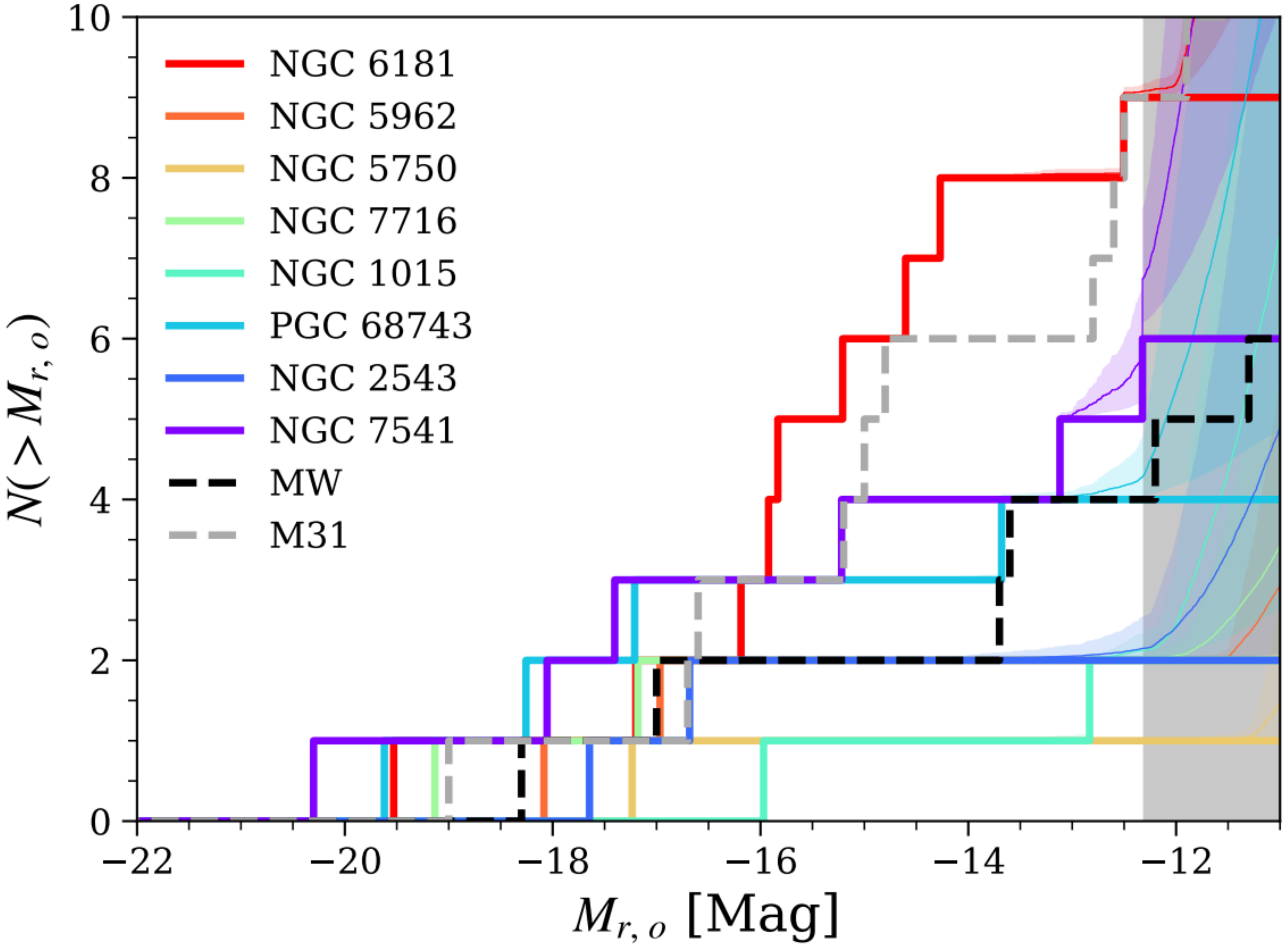


Boylan-Kolchin et al. 2015

Weisz & Boylan-Kolchin 2017

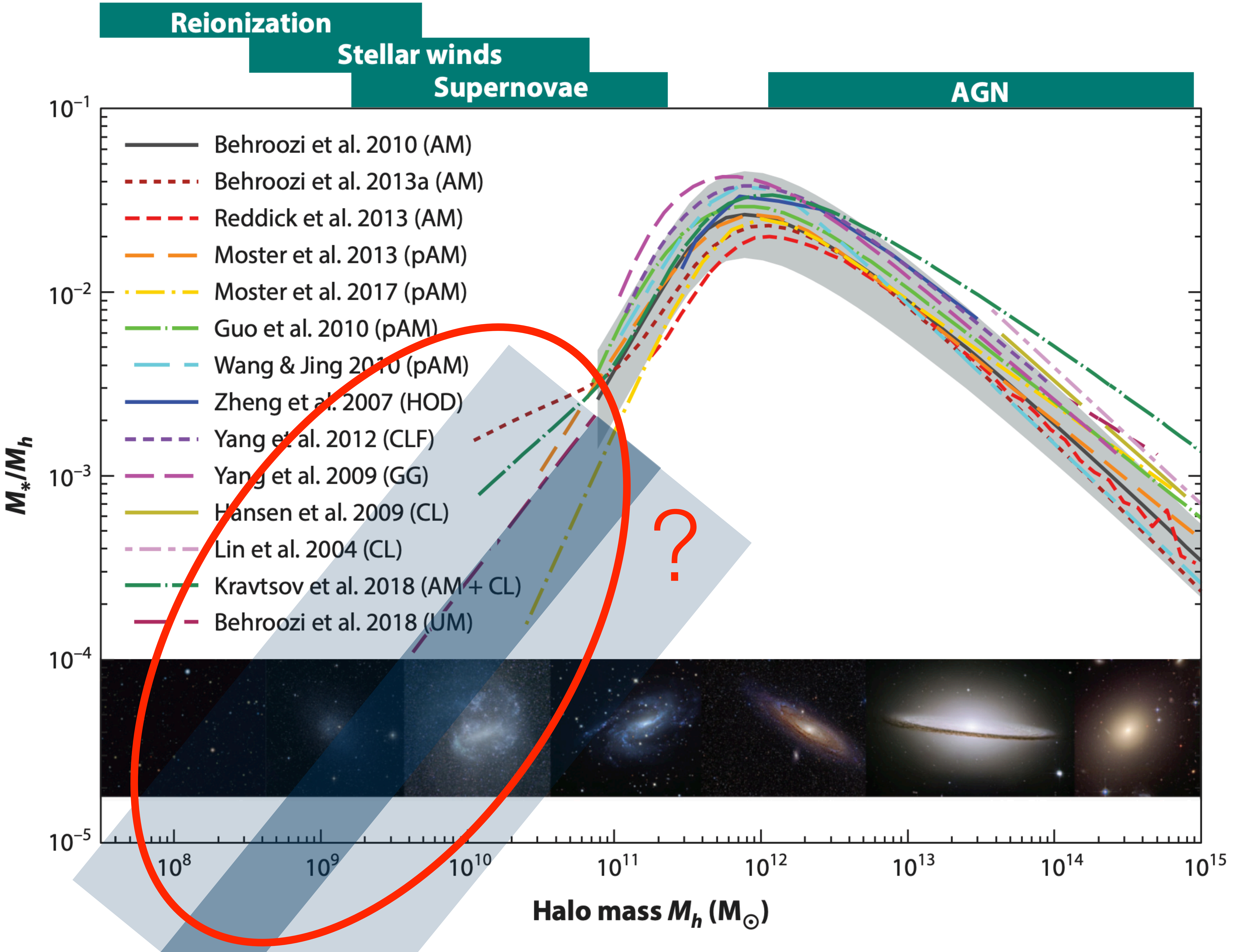
Interplay with Local Satellite Surveys

- Surveys of Milky Way-like satellite systems in the Local Universe are rapidly progressing
- The Satellites Around Galactic Analogs (SAGA) survey targets MW analogs within 20-40 Mpc
- MW satellite luminosity function is typical among SAGA, LV systems
- Detailed work to jointly fit the MW and other systems in progress!



Geha et al. 2017

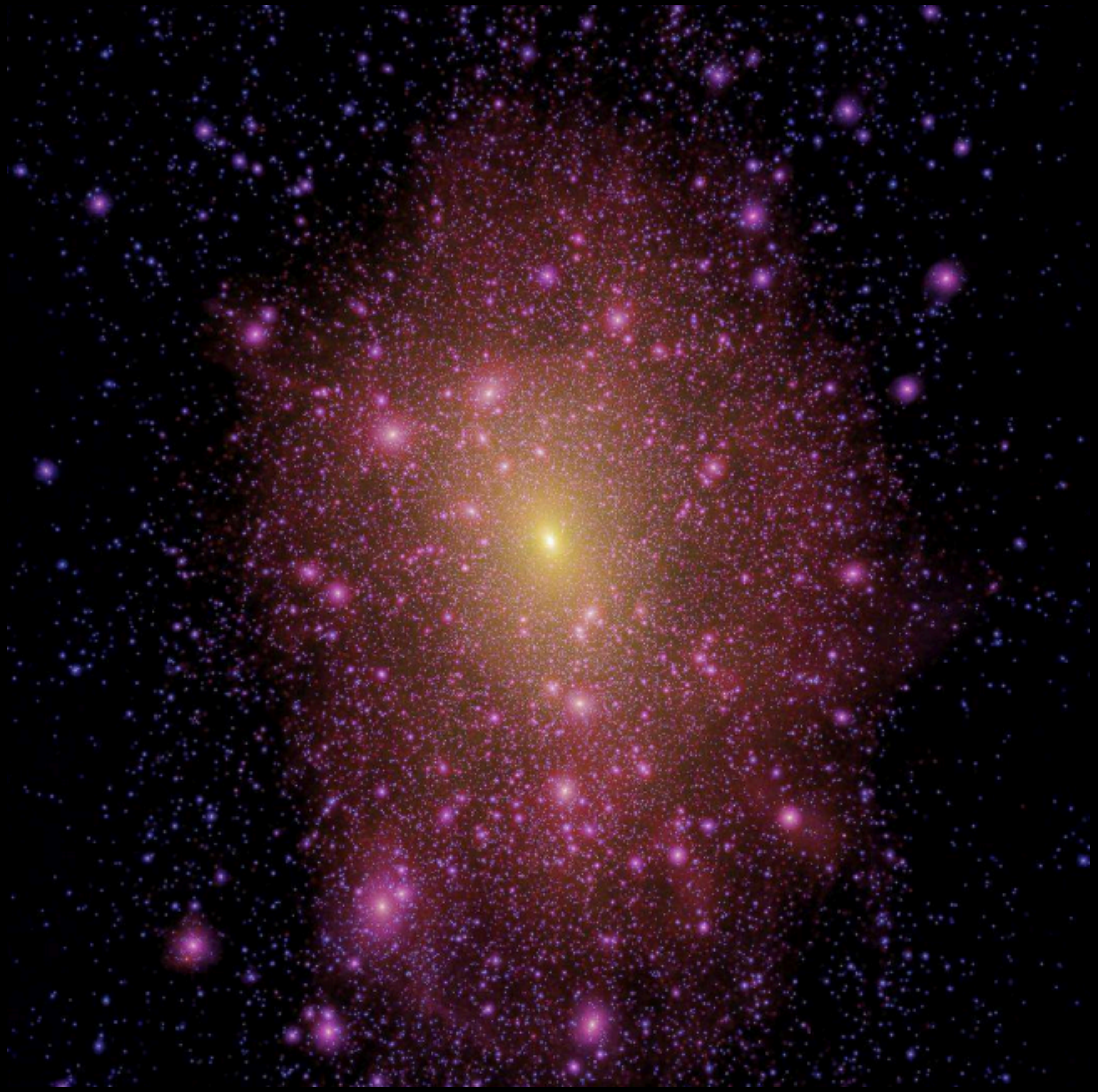
The Faint-End Galaxy–Halo Connection



Wechsler & Tinker 2018

Dark Matter Constraints

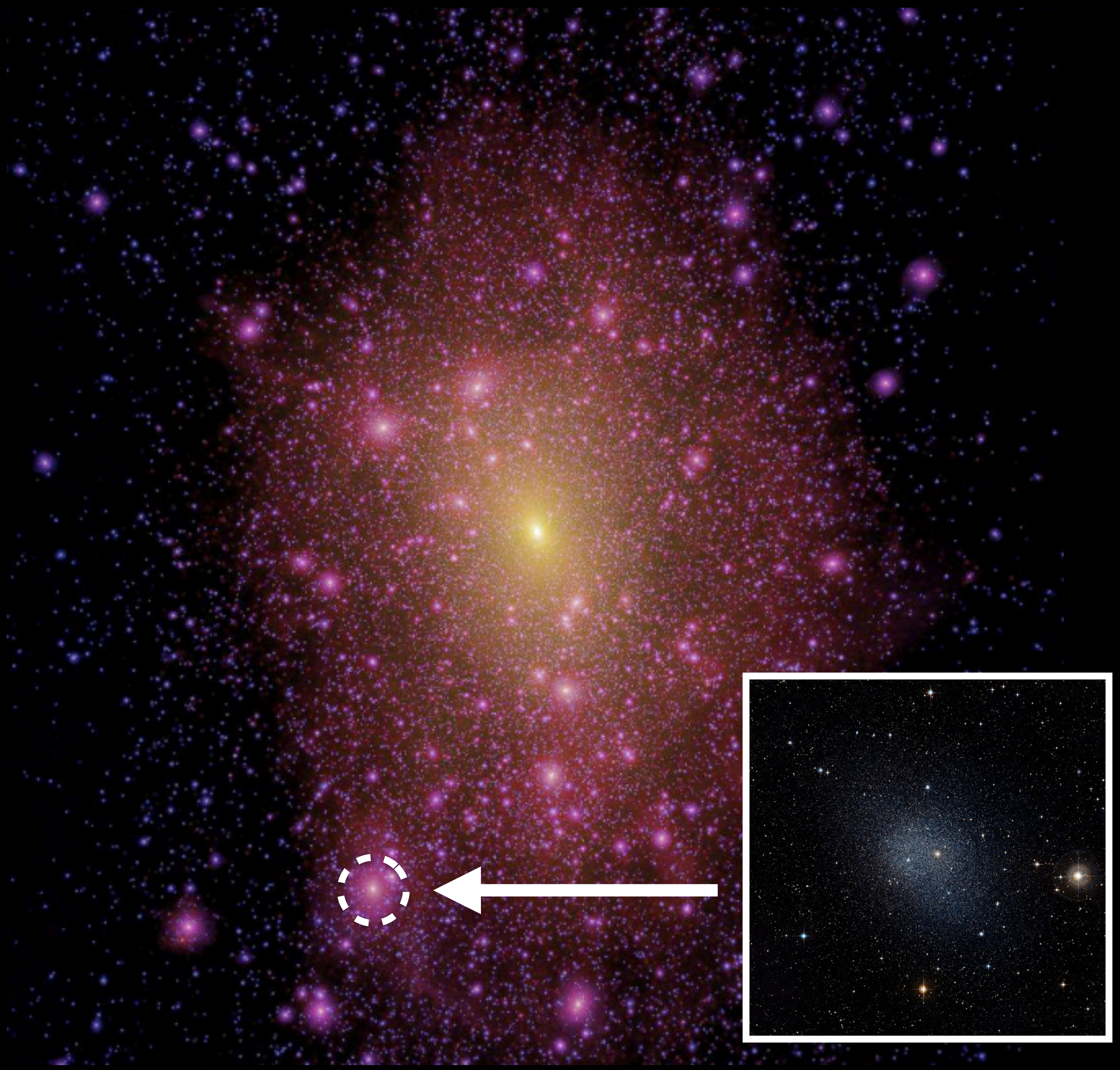
CDM

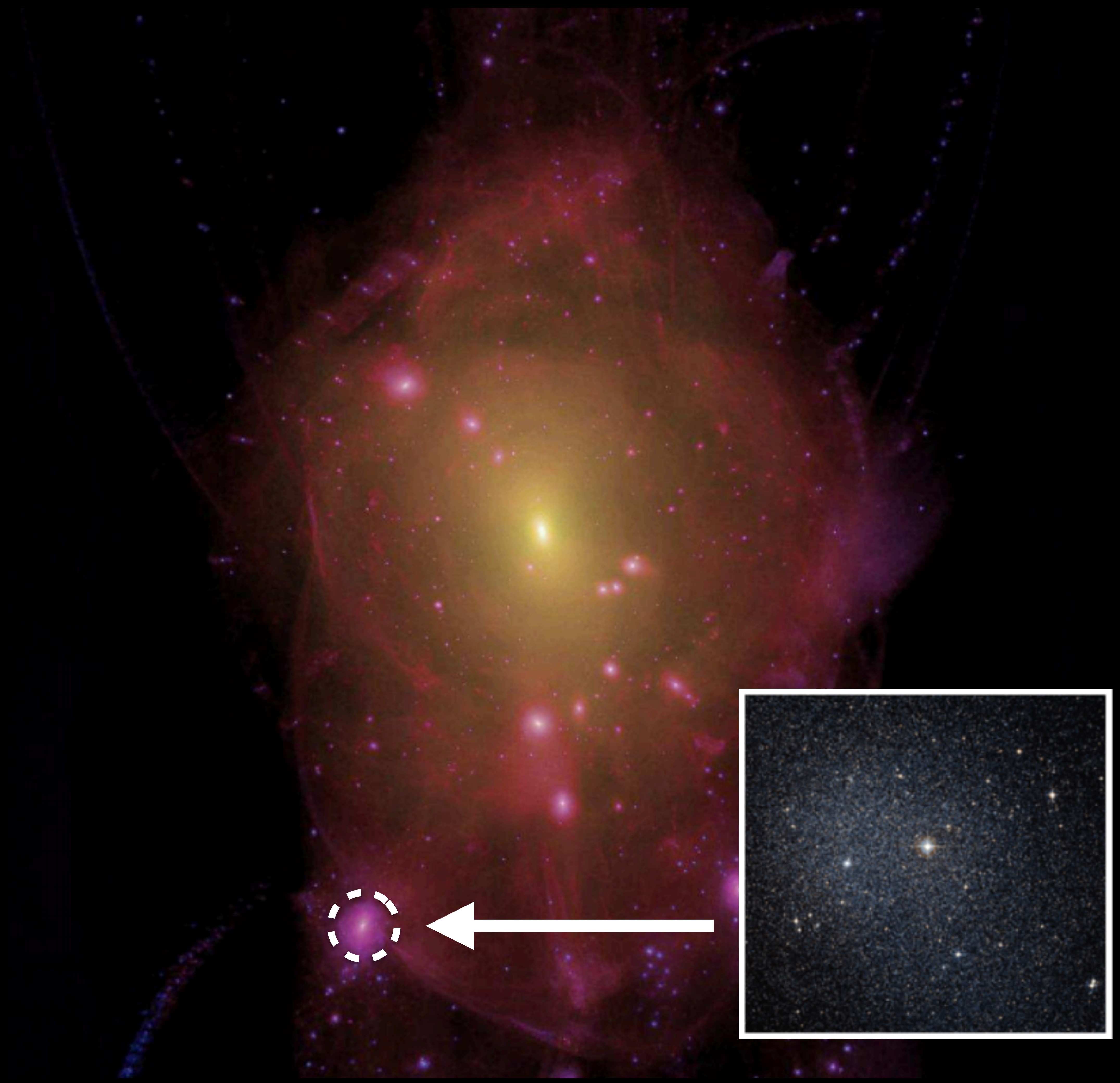


WDM

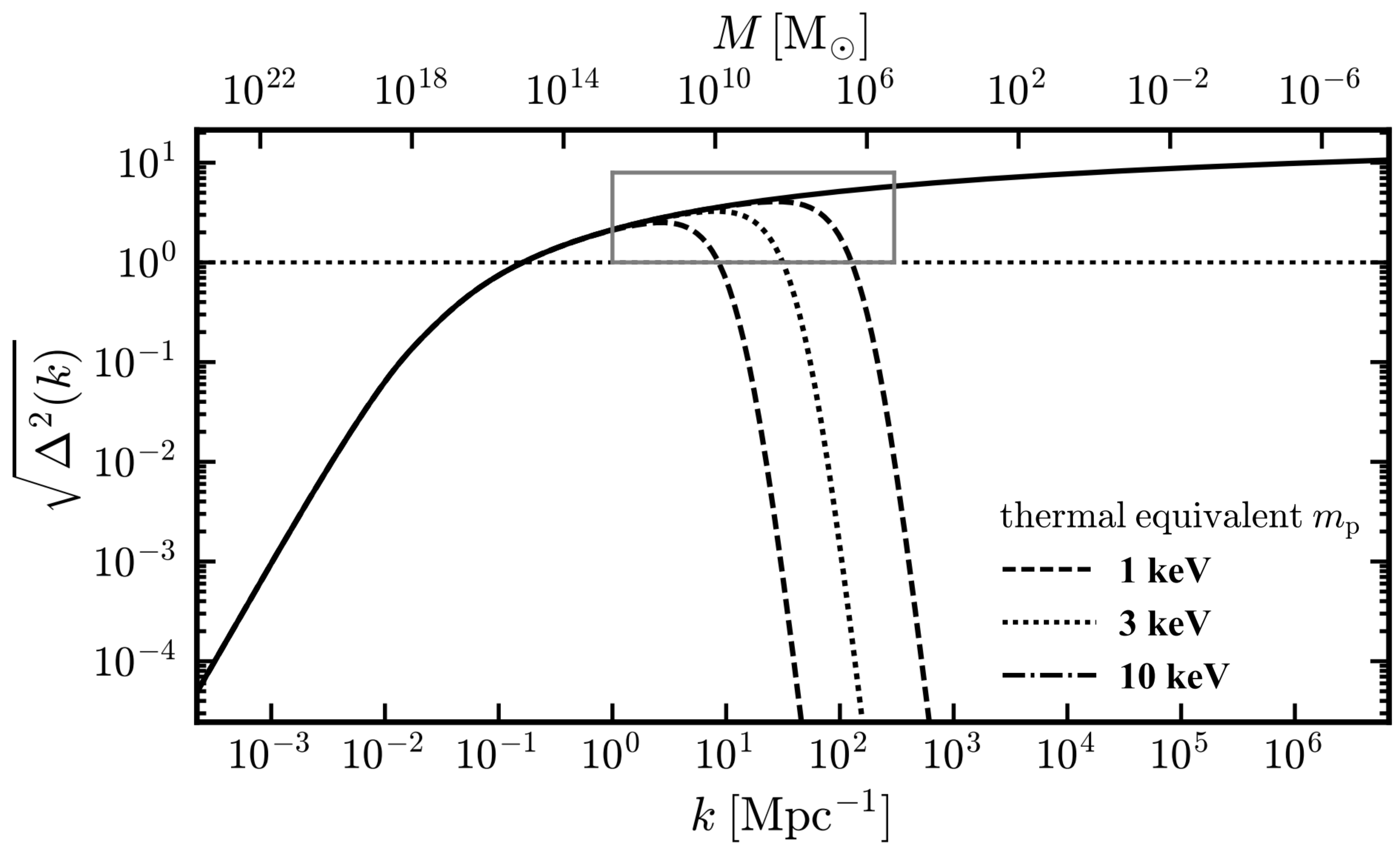
Lovell et al. 2011

Dark Matter Constraints

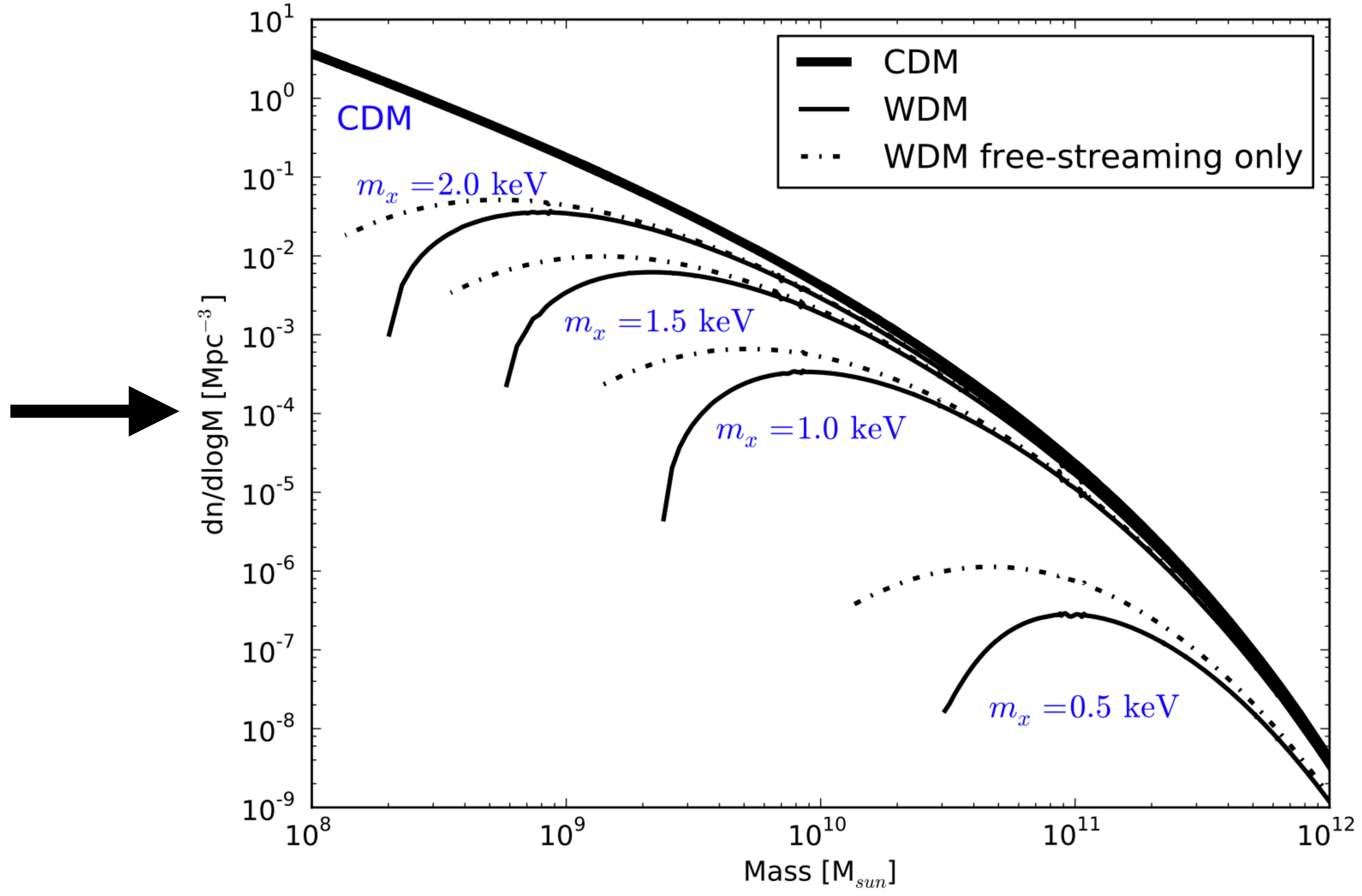




DM Physics and the Minimum Halo Mass



Bullock & Boylan-Kolchin 2017

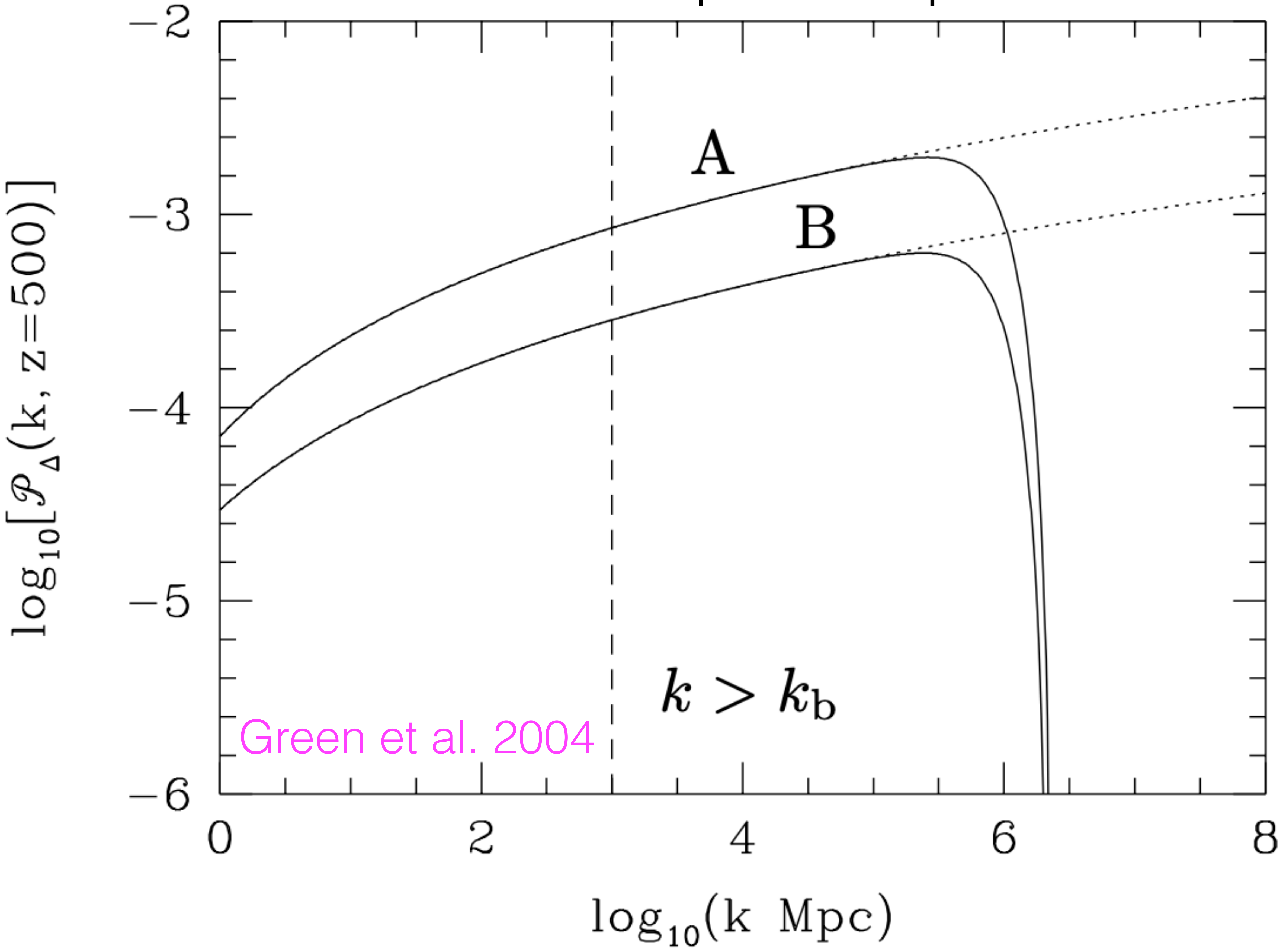


Pacucci et al. 2013

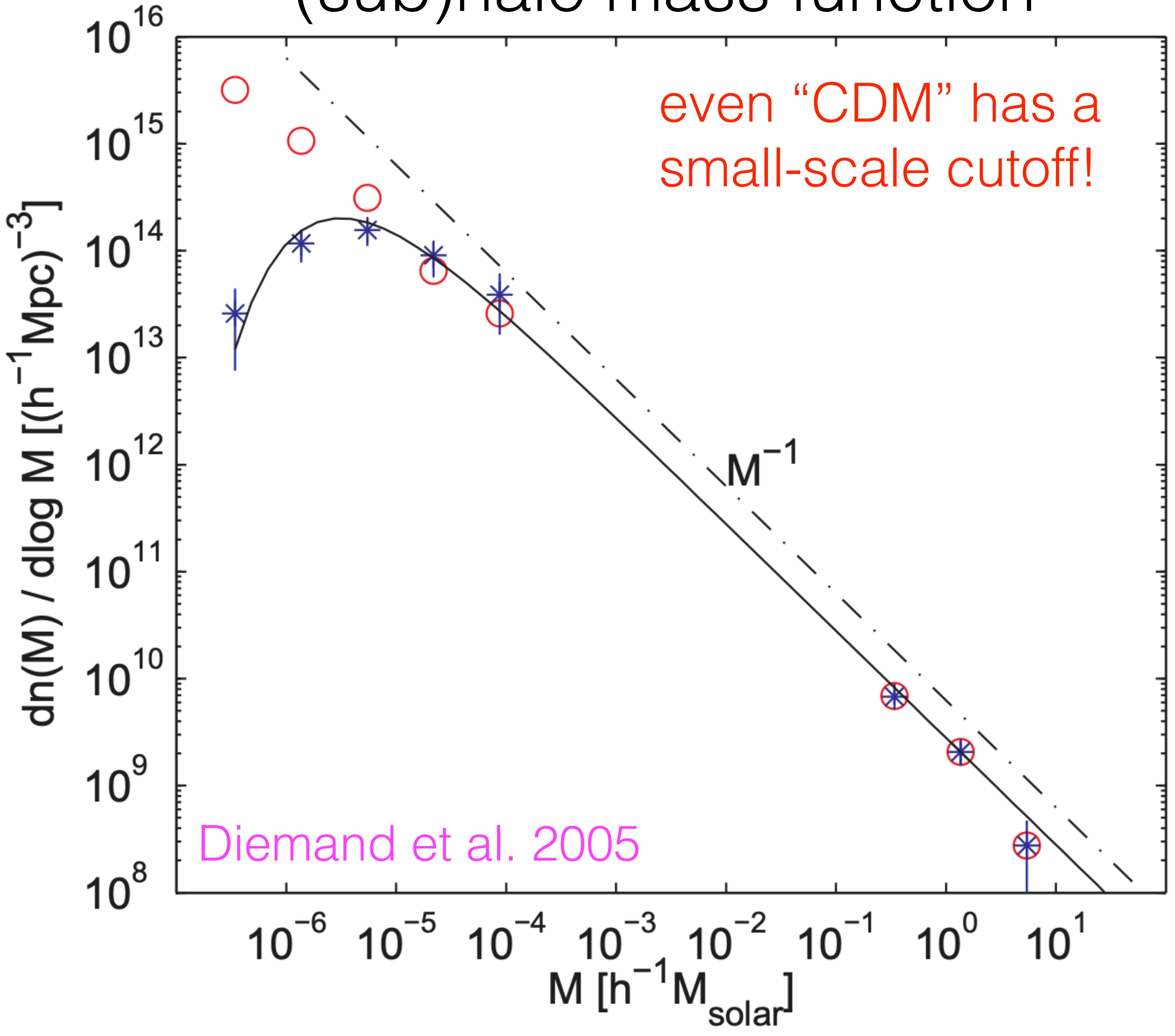
free-streaming due to large primordial velocity dispersion → cutoff in abundance of low-mass halos

DM Physics and the Minimum Halo Mass

linear matter power spectrum

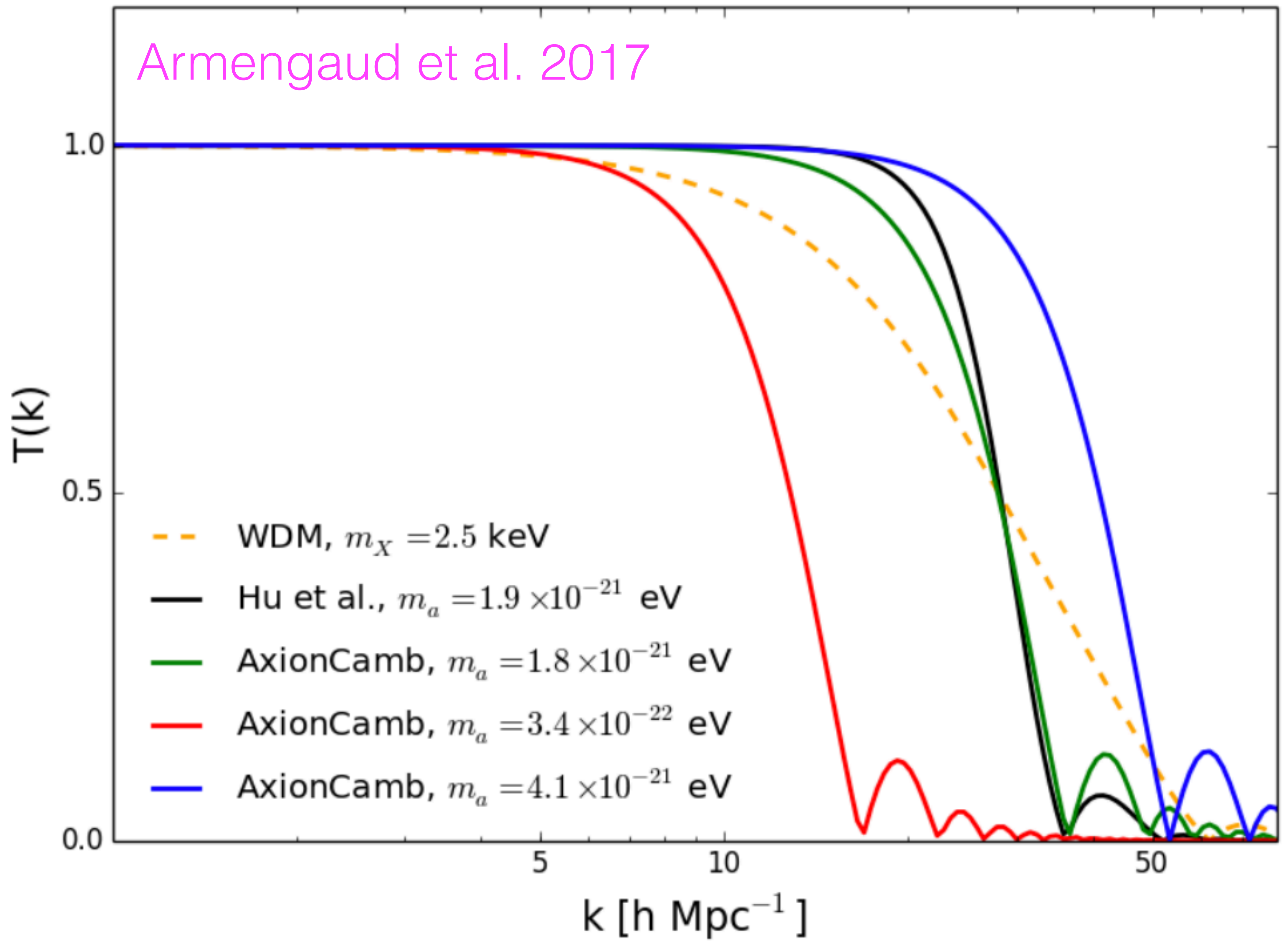


(sub)halo mass function



collisional damping due to DM–SM interactions → cutoff in abundance of low-mass halos

DM Physics and the Minimum Halo Mass



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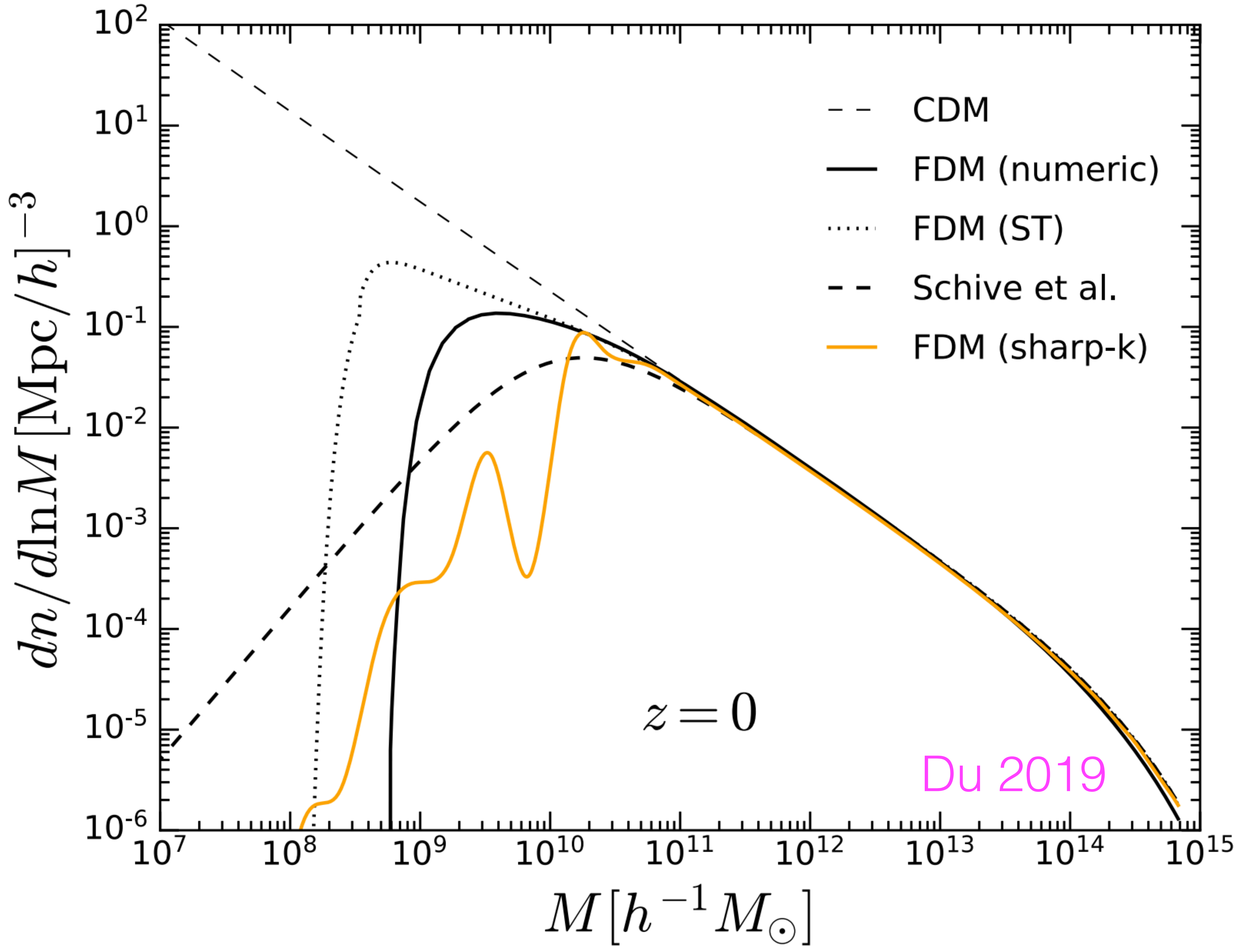
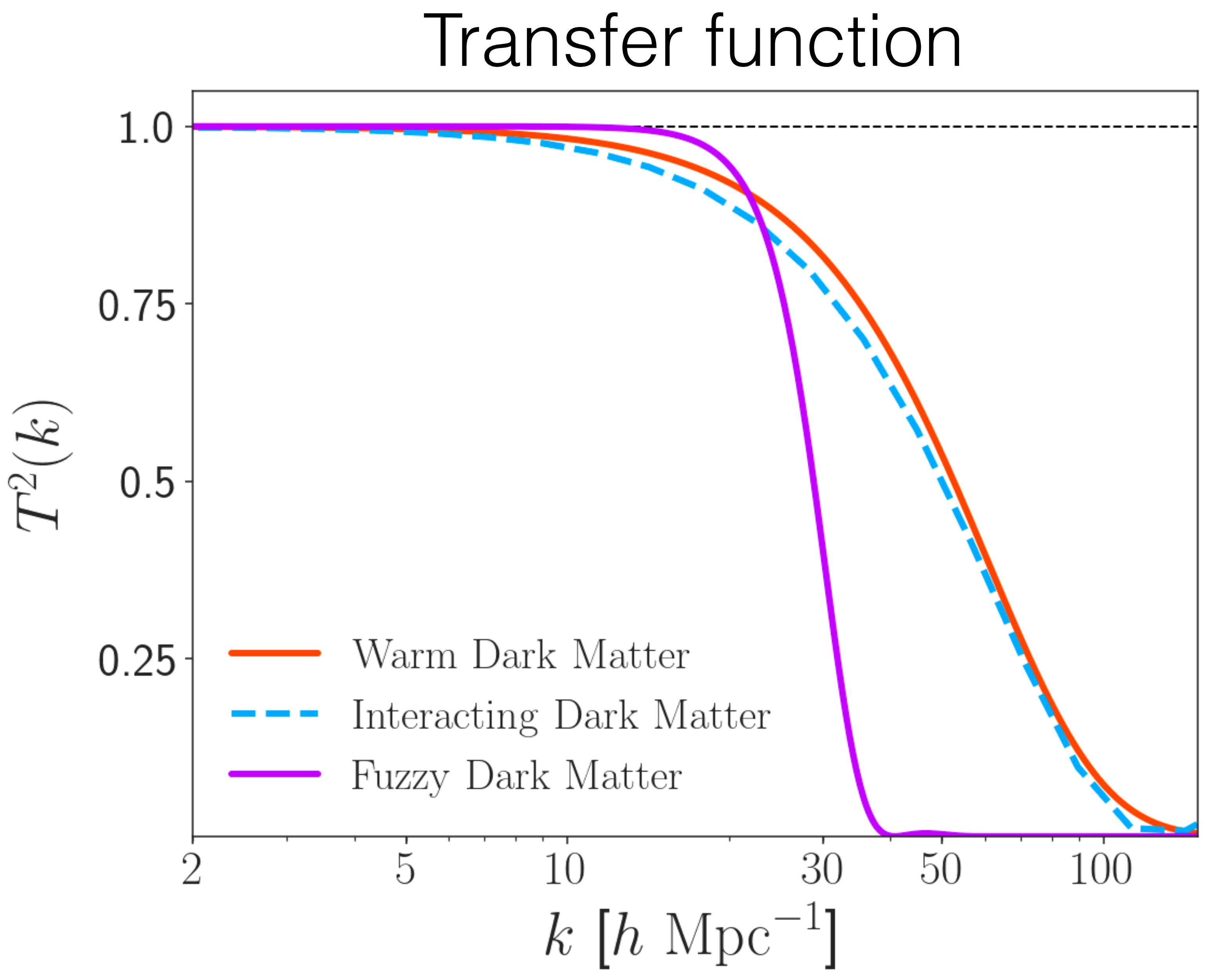
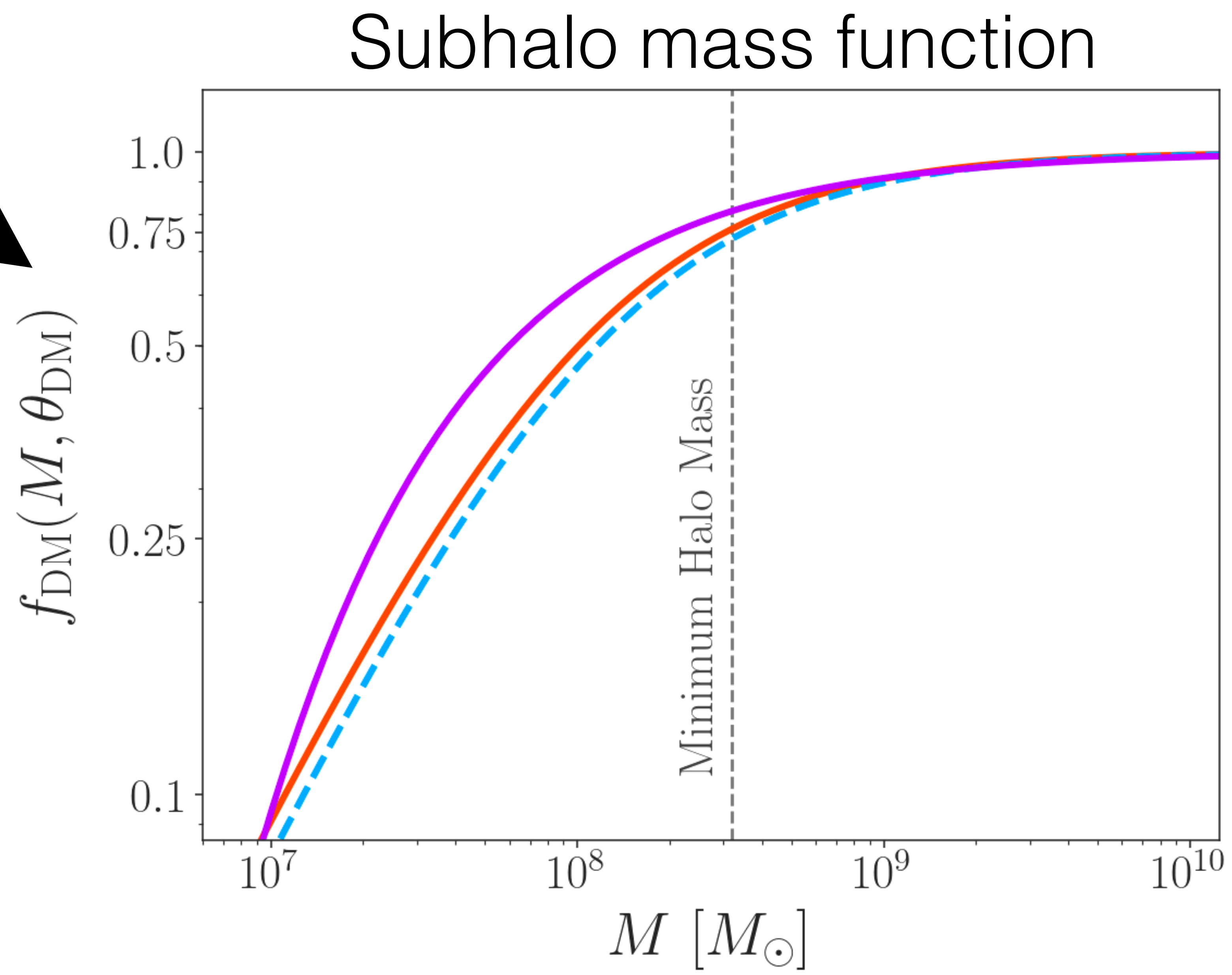
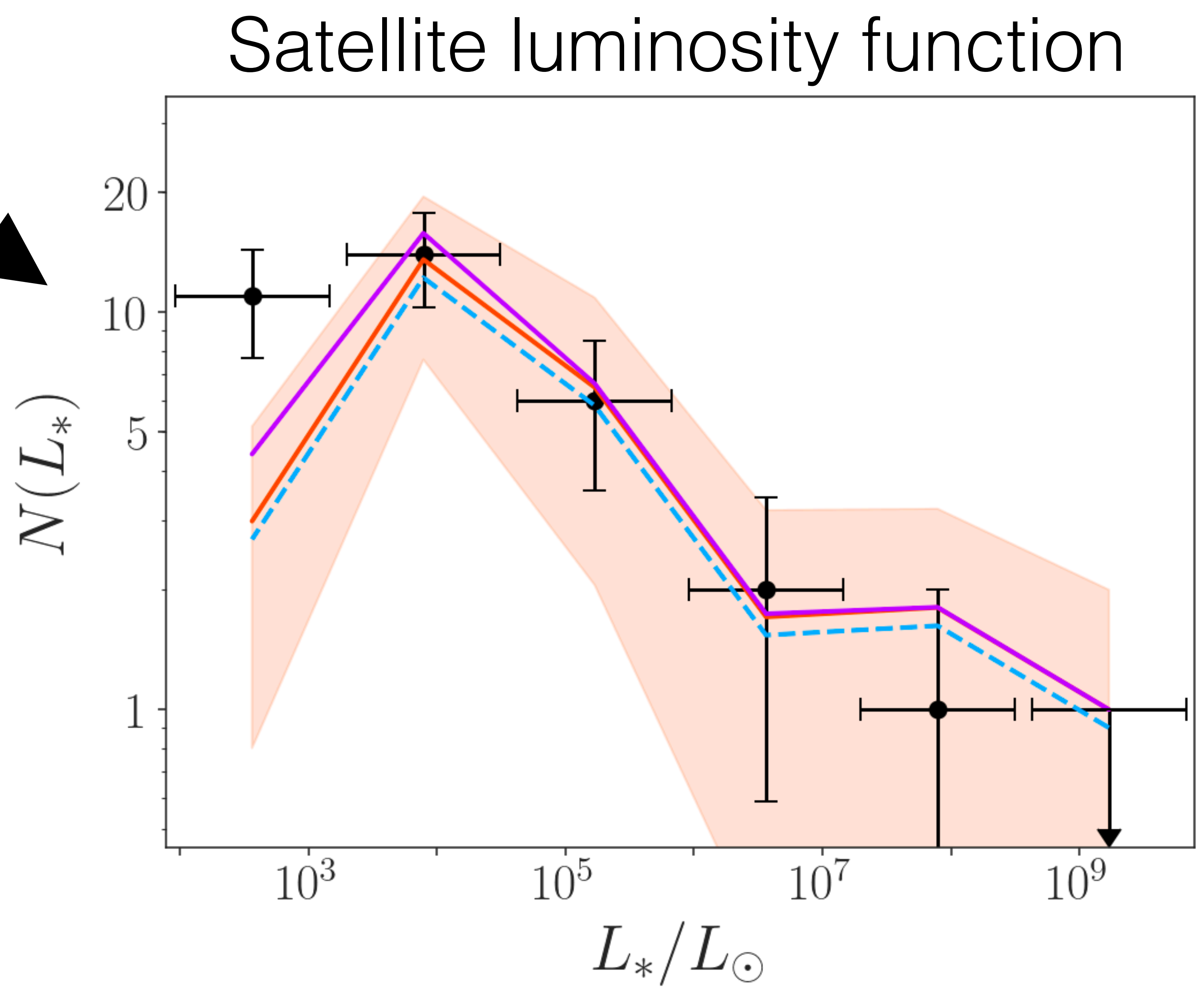


interference due to macroscopic de Broglie wavelength → cutoff in abundance of low-mass halos



Fit the Milky Way satellite population with subhalo mass function suppression, marginalizing over galaxy–halo connection uncertainties and Milky Way halo properties:





Current satellite observations are sensitive to **~25%** suppression in subhalo abundance relative to CDM

EN & Drlica-Wagner et al. 2020

Dark Matter Constraints

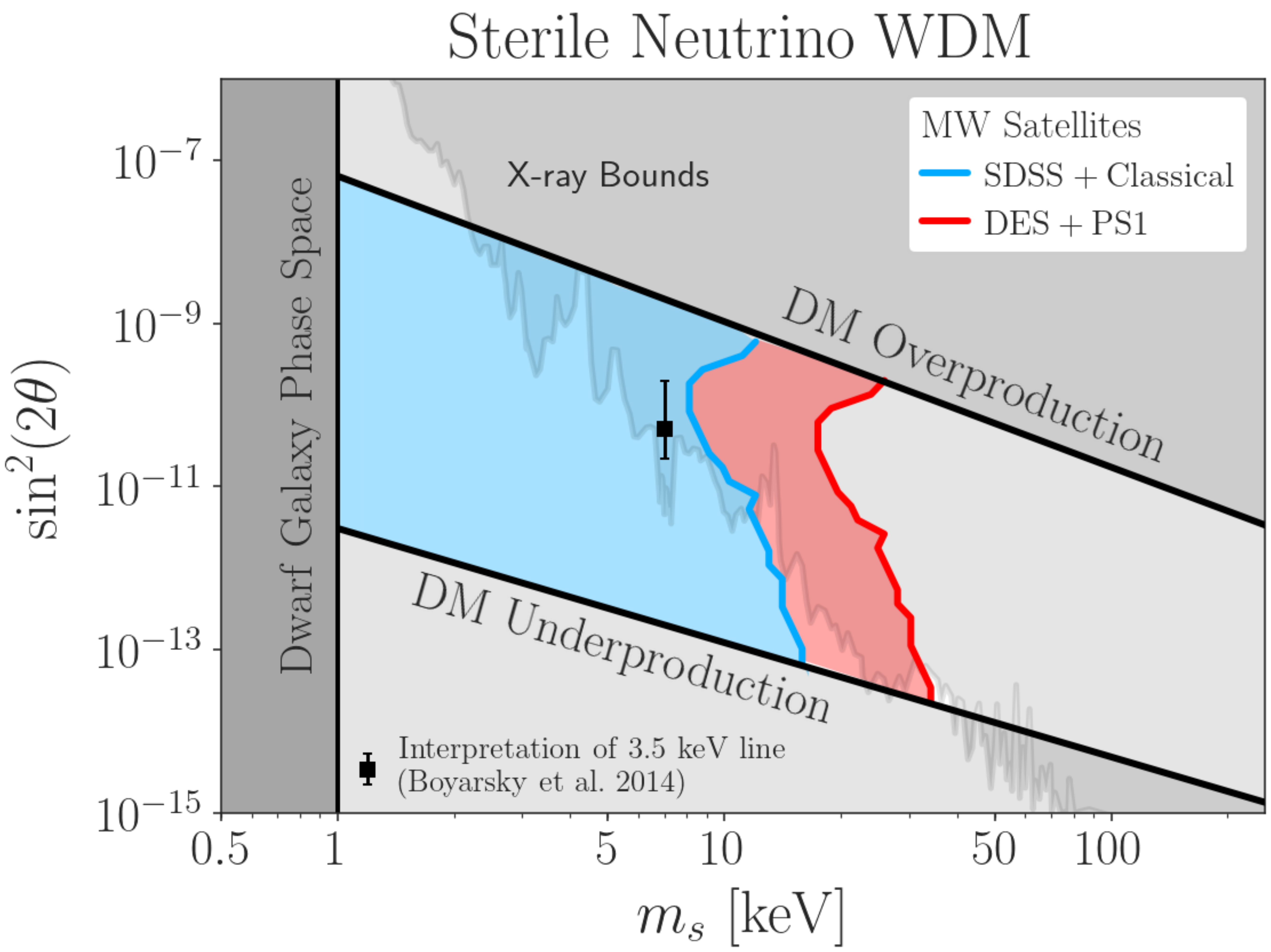
- Our Milky Way satellite analysis yields the strongest astrophysical constraints to date on a variety of dark matter properties and particle models
- For thermal relic WDM: $\boldsymbol{m_{\mathrm{WDM}} > 6.5\ \mathrm{keV}}$ (95% confidence), comparable to limits from the Lyman-alpha forest, strong gravitational lensing, and stellar stream perturbations (with largely distinct systematics)
- This is the opposite of searching under a lamppost!

Dark Matter Paradigm	Parameter	Constraint	Derived Property	Constraint
Warm Dark Matter	Thermal Relic Mass	$m_{\mathrm{WDM}} > 6.5\,\mathrm{keV}$	Free-streaming Length	$\lambda_{\mathrm{fs}} \lesssim 10\,h^{-1}\,\mathrm{kpc}$
Interacting Dark Matter	Velocity-independent DM–Proton Cross Section	$\sigma_0 < 8.8 \times 10^{-29}\,\mathrm{cm}^2$	DM–Proton Coupling	$c_p \lesssim (0.3\,\mathrm{GeV})^{-2}$
Fuzzy Dark Matter	Particle Mass	$m_\phi > 2.9 \times 10^{-21}\,\mathrm{eV}$	de Broglie Wavelength	$\lambda_{\mathrm{dB}} \lesssim 0.5\,\mathrm{kpc}$

EN & Drlica-Wagner et al. 2020

Warm Dark Matter Constraints

- Our analysis excludes nearly all remaining parameter space for resonantly-produced sterile neutrino dark matter
- The sterile neutrino interpretation of the 3.5 keV X-ray line is ruled out at >> 99% confidence
- Interpretation: the dark matter free-streaming length must be smaller than the sizes of the halos that host ultra-faint dwarf galaxies



EN & Drlica-Wagner et al. 2020

Interacting Dark Matter Constraints

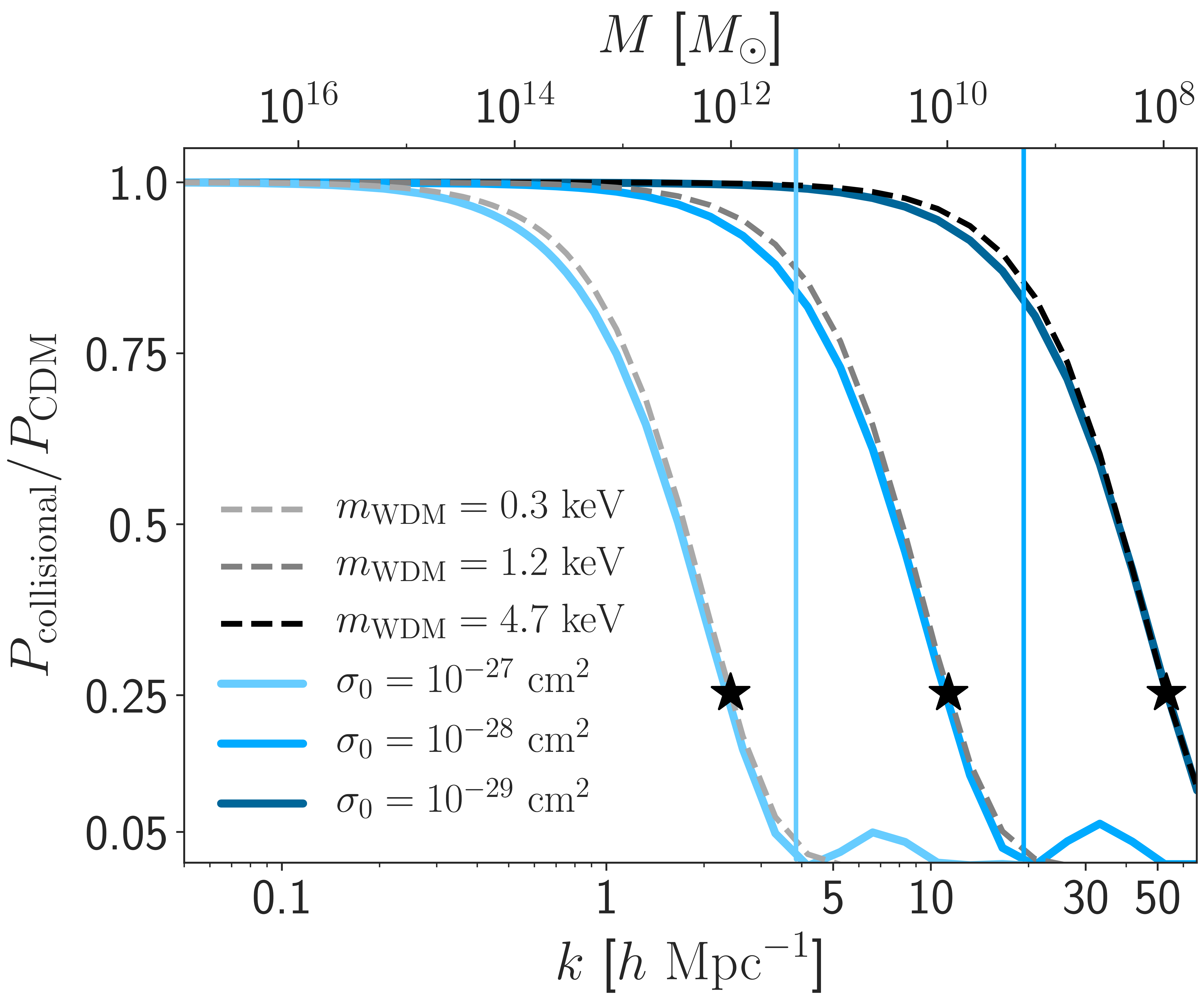
- Collisional damping due to DM–baryon scattering at early times suppresses power on small scales

- Mass of the smallest halo allowed to form corresponds to the size of the horizon when

$$R_\chi \sim aH$$

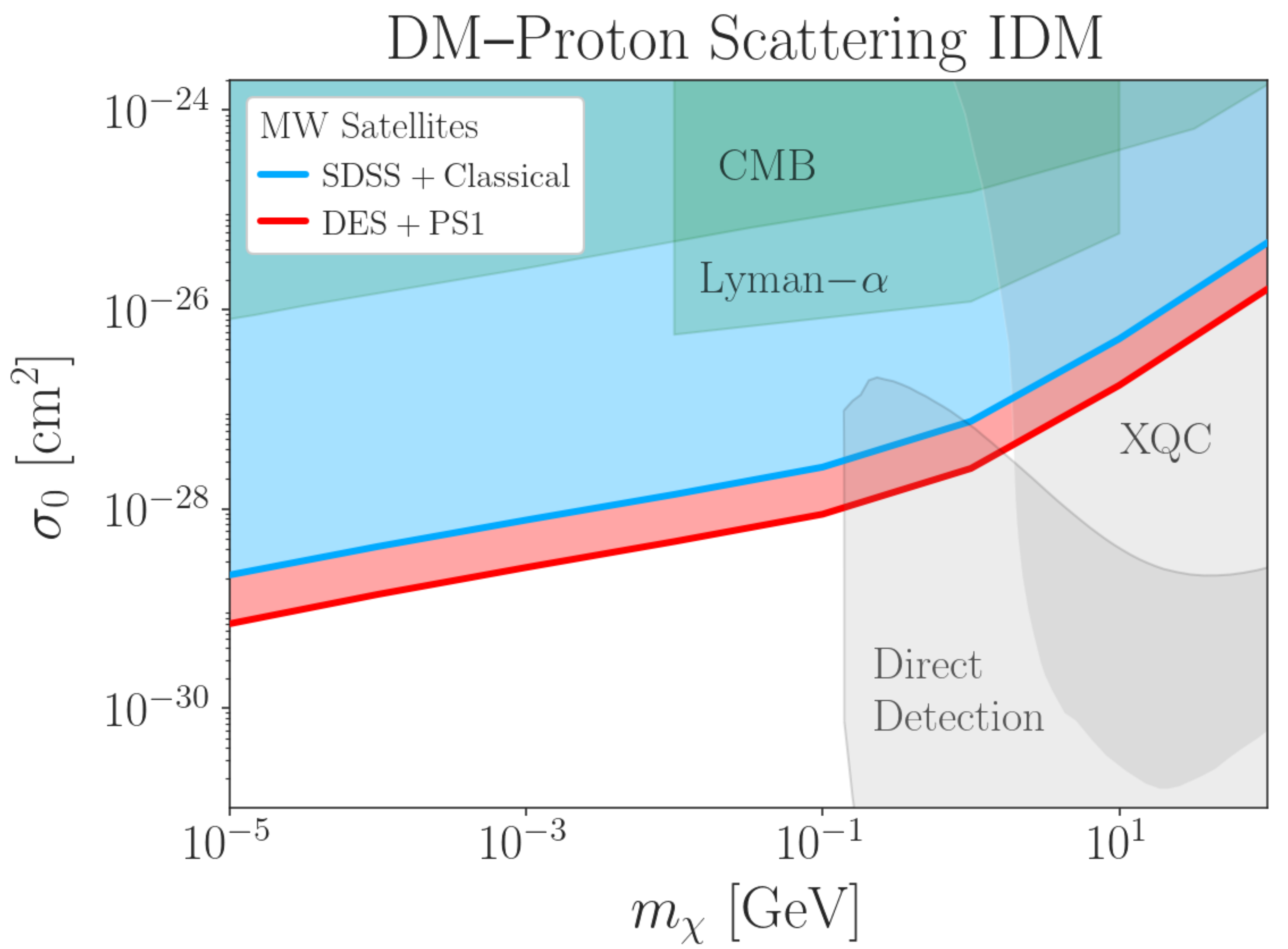
Momentum transfer rate → ← Hubble rate

- Minimum observed halo mass yields analytic cross section limit



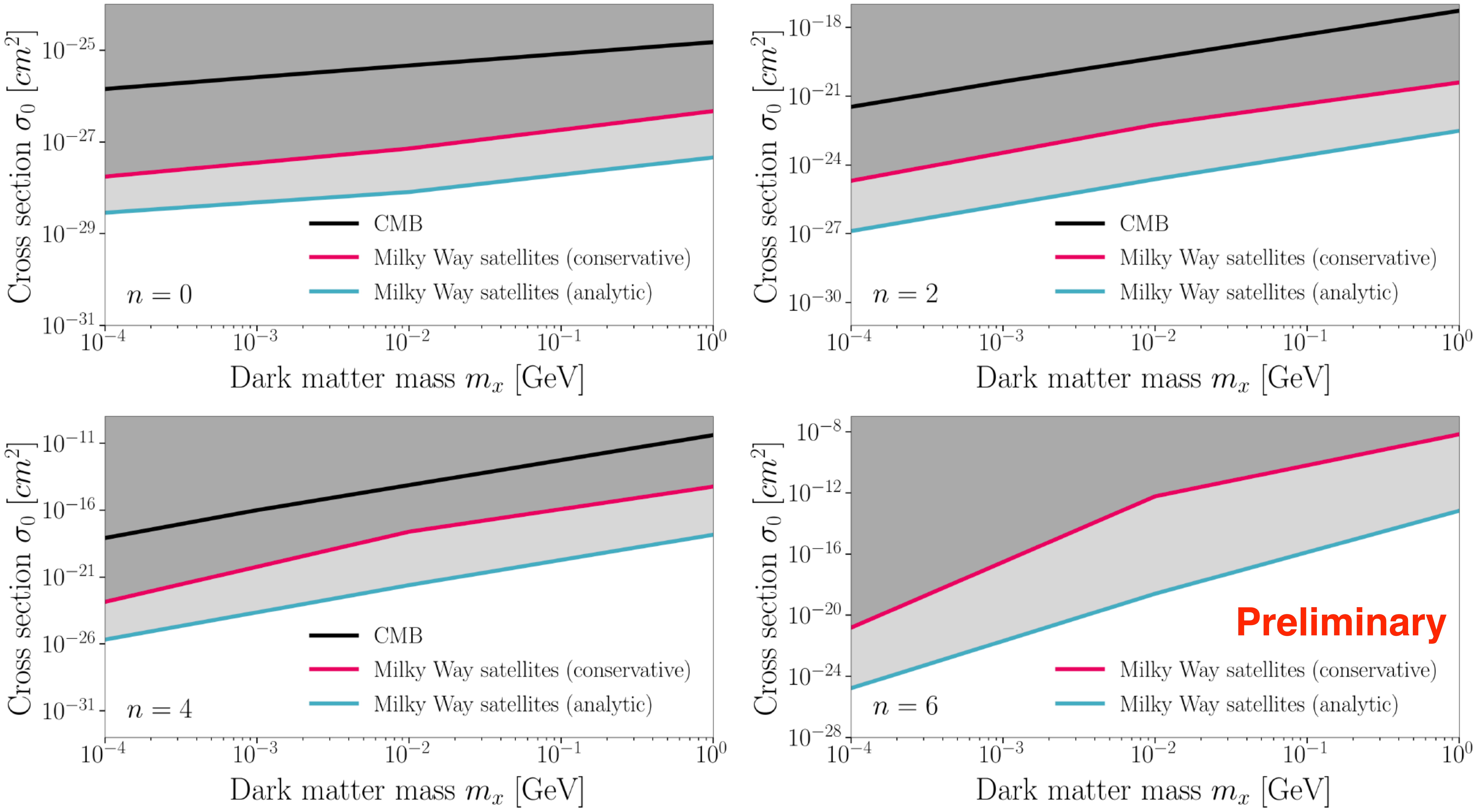
EN, V. Gluscevic, K. Boddy, R. Wechsler 2019

Interacting Dark Matter Constraints



EN & Drlica-Wagner et al. 2020

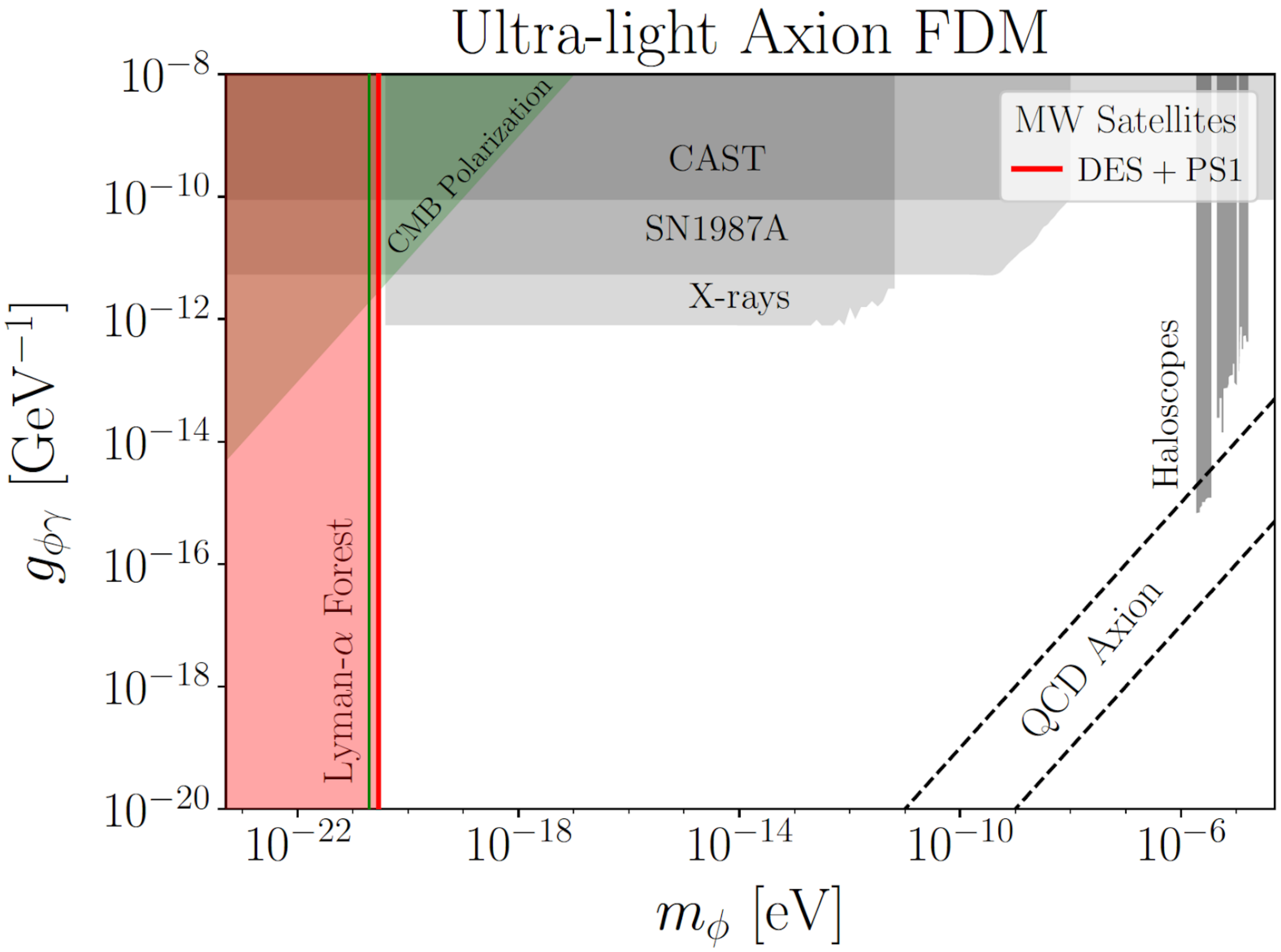
Interacting Dark Matter Constraints



K. Maamari et al. in prep

Fuzzy Dark Matter Constraints

- Our analysis robustly rules out fuzzy dark matter masses below 10^{-21} eV, with very conservative modeling assumptions
- Can be interpreted as a lower limit on the ultra-light axion mass assuming negligible self and Standard Model couplings
- Implies that the DM de Broglie wavelength is smaller than the sizes of ultra-faint dwarfs

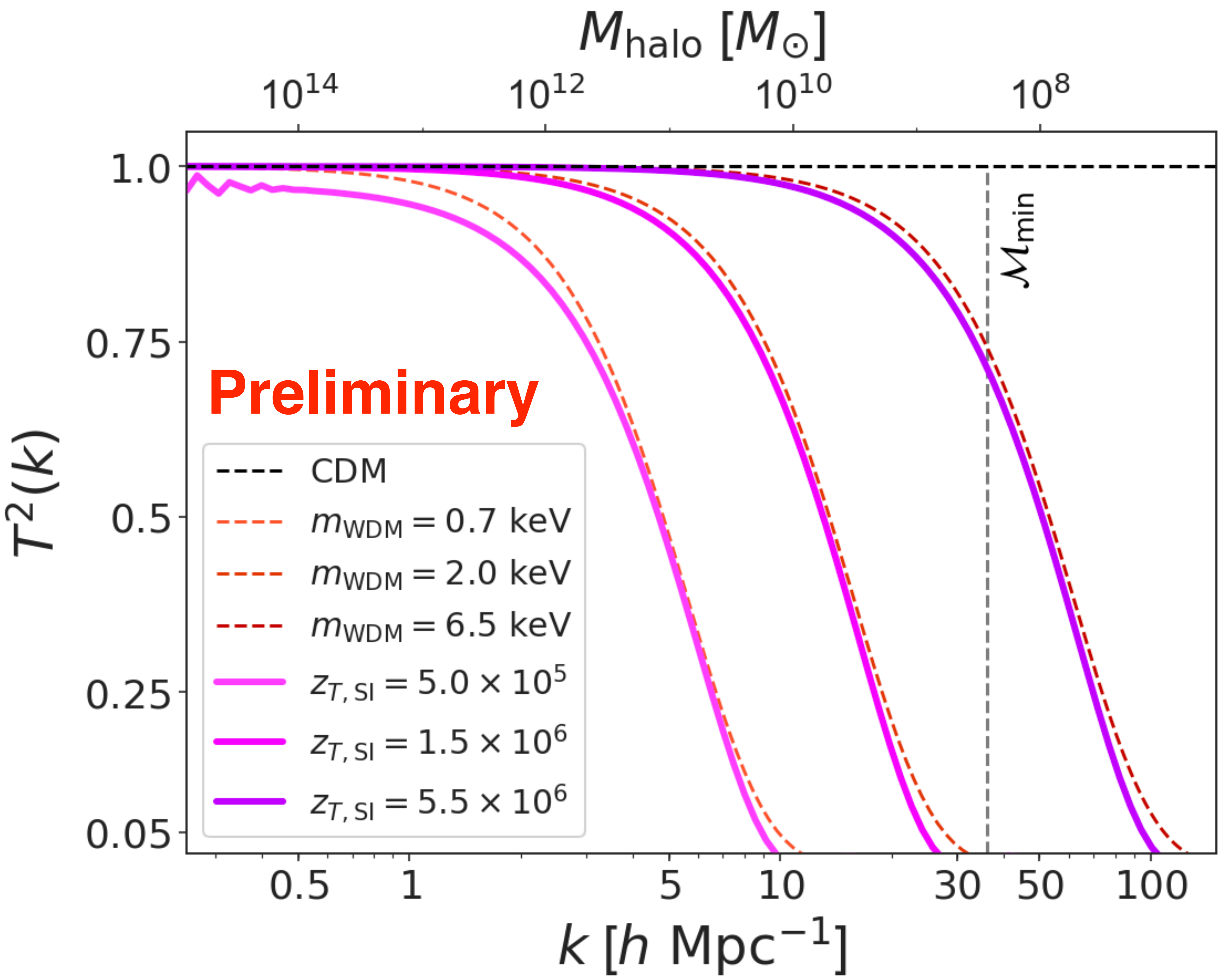


EN & Drlica-Wagner et al. 2020

Late-Forming Dark Matter Constraints

- Dark radiation that decays to DM after BBN ("late-forming DM") suppresses power on small scales
- Cutoff in halo abundance is set by the LFDM transition redshift
- MW satellites yields an order-of-magnitude improvement on the transition redshift lower bound:

$$z_{\mathrm{LFDM}} > 5.5 \times 10^{6}\ (95\%\ \mathrm{C.L.})$$

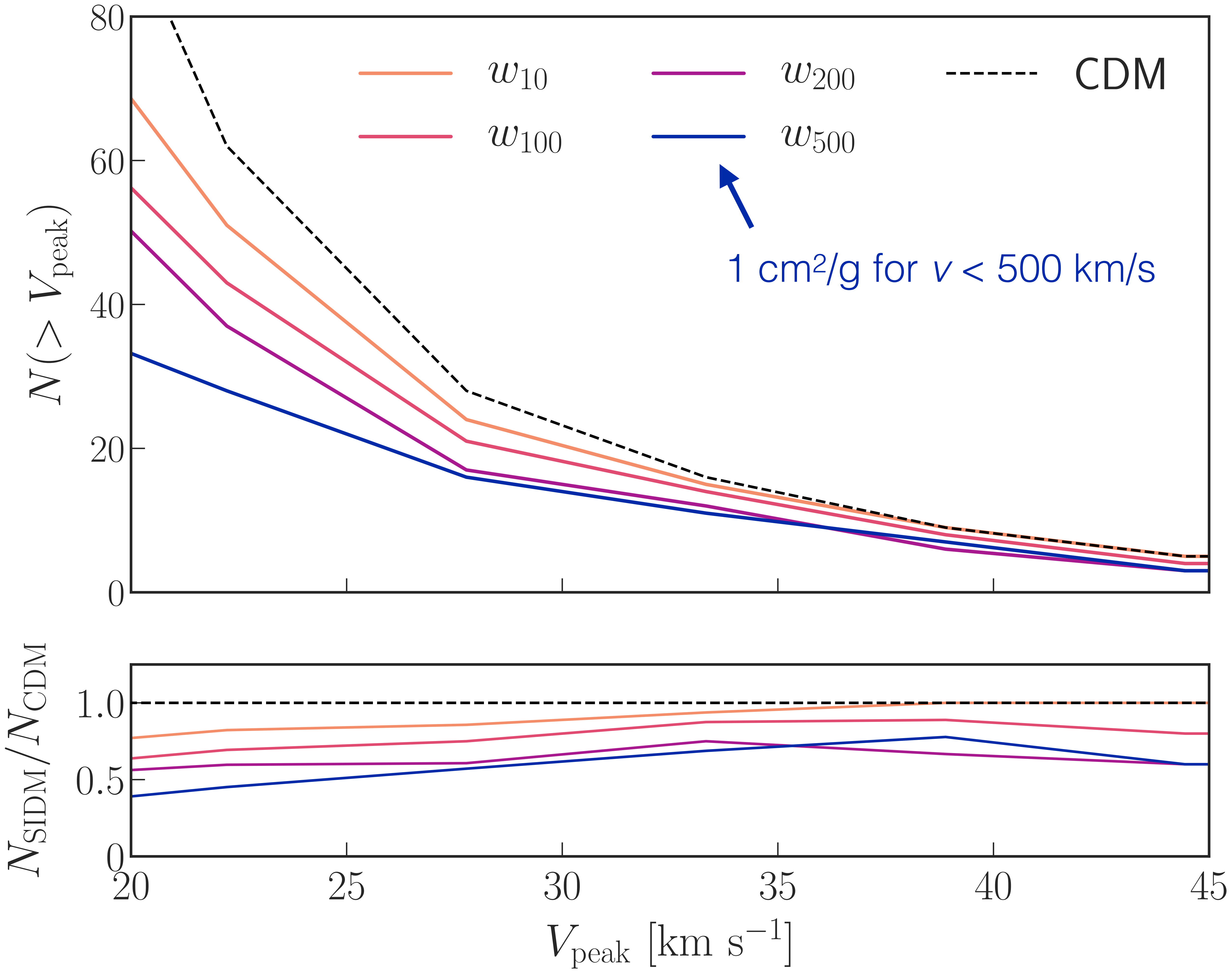


S. Das & EN in prep.

Constraining Late-time DM Physics

- Our framework also constrains DM properties that suppress subhalo abundances at late times (e.g. DM self-interactions, decays)

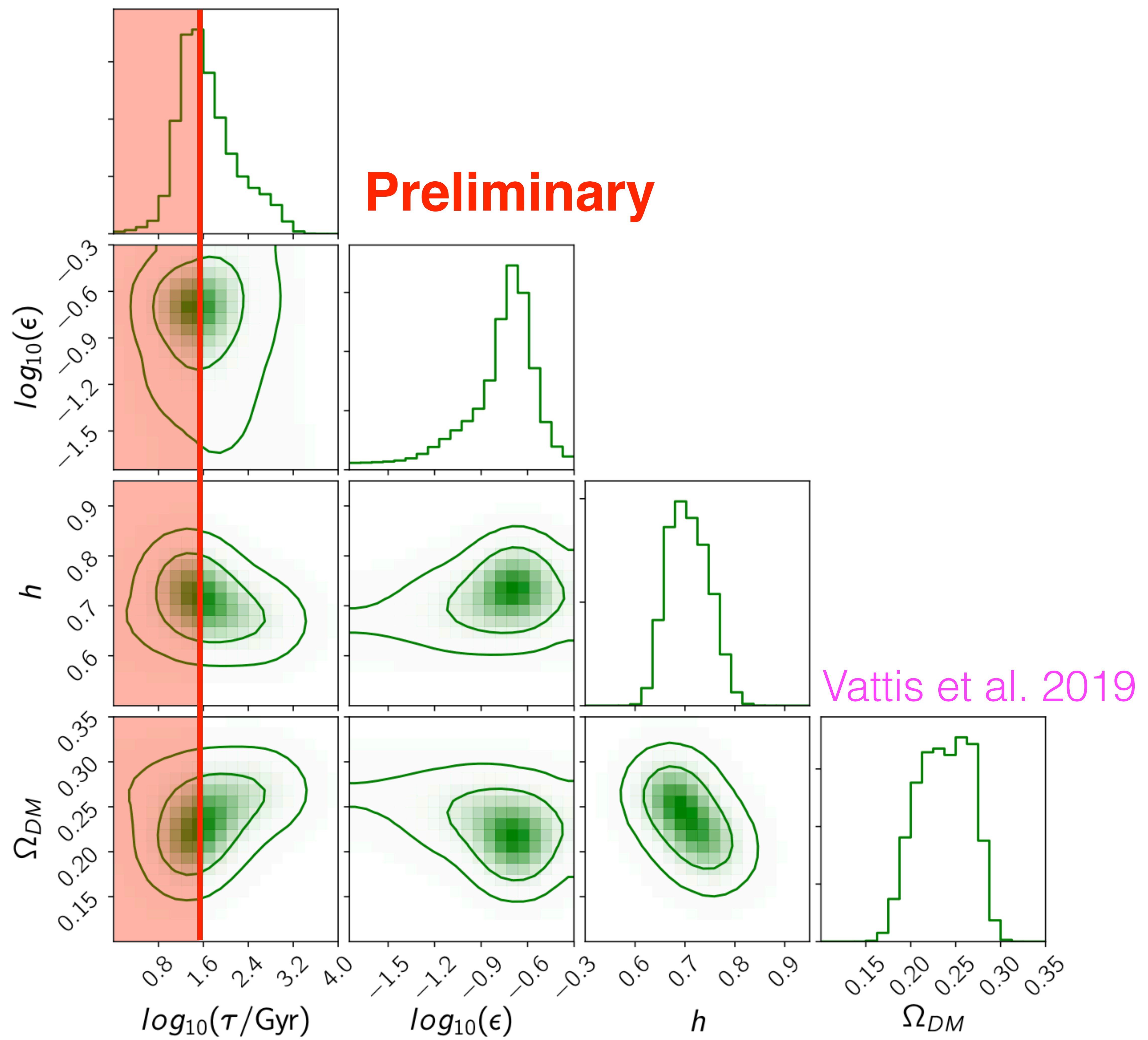
- Self-interactions: Sensitivity to SIDM cross sections of ~1 cm²/g at low velocity scales (stay tuned!)



EN, A. Banerjee, S. Adhikari, Y.-Y. Mao, R. Wechsler 2020

Constraining Late-time DM Physics

- Our framework also constrains DM properties that suppress subhalo abundances at late times (e.g. DM self-interactions, decays)
- Decays: Sensitivity to decaying DM with a lifetime of ~10 Gyr, which has been claimed to alleviate the Hubble tension (stay tuned!)



S. Mau et al. in prep

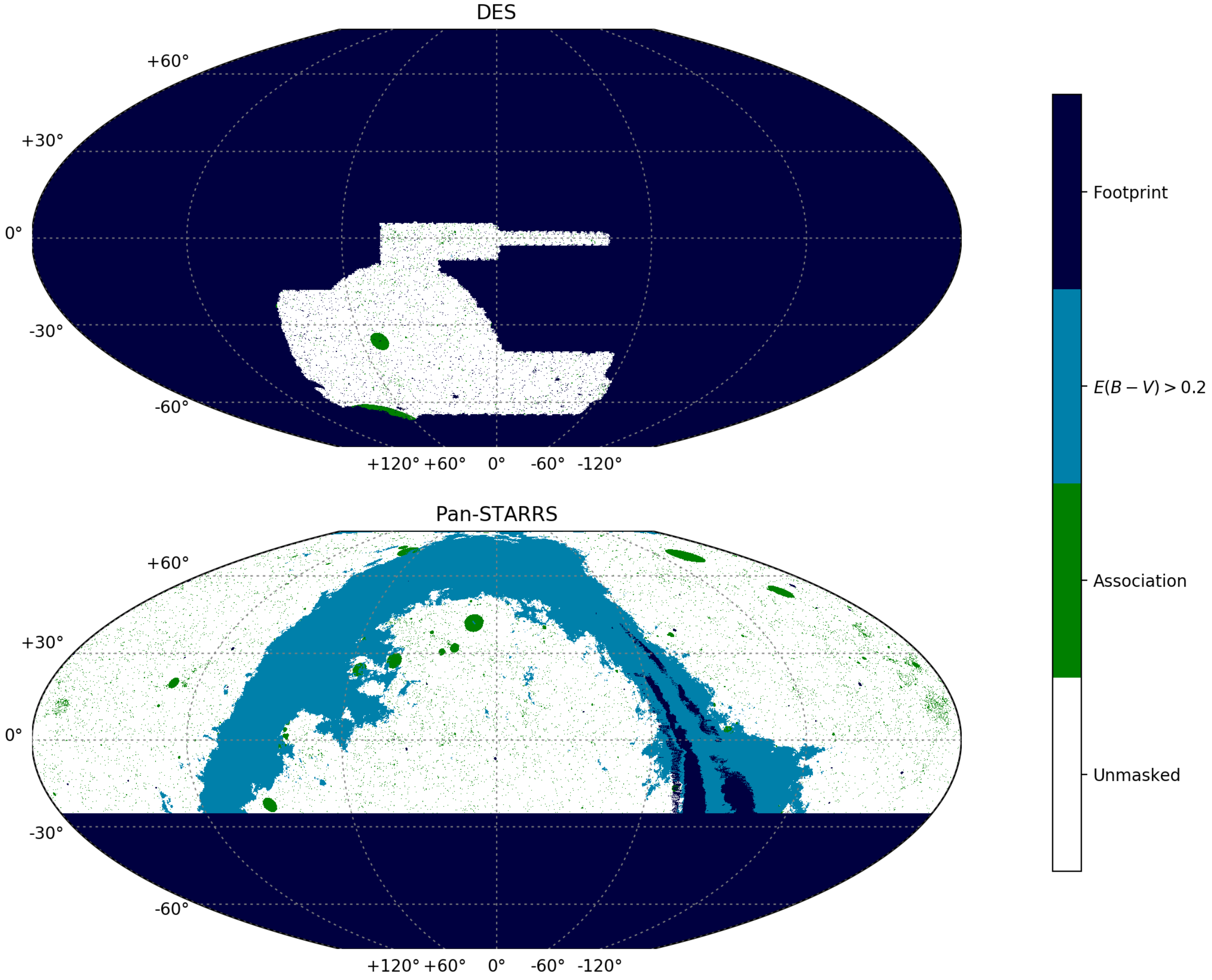
Outlook

- The lack of a cutoff in the abundance of observable ultra-faint galaxies down to halo masses of $\sim 10^8\ M_\odot$ yields stringent constraints on the galaxy—halo connection and dark matter physics
- These constraints will **continue to improve** with advances in (currently conservative) galaxy–halo modeling and LSST satellite discoveries
- Our analysis informs a variety of DM properties (free-streaming scale, de Broglie wavelength, coupling to the Standard Model) and particle models (sterile neutrinos, generalized WIMPs, ultra-light axions)
- Joint modeling of small-scale structure probes is an important area for future work, starting with satellites, streams, and strong lenses

Thanks!

Susmita Adhikari, Arka Banerjee, Keith Bechtol, Kimberly Boddy,
Subinoy Das, Alex Drlica-Wagner, Vera Gluscevic, Greg Green,
Yao-Yuan Mao, Sidney Mau, Mitch McNanna, Risa Wechsler

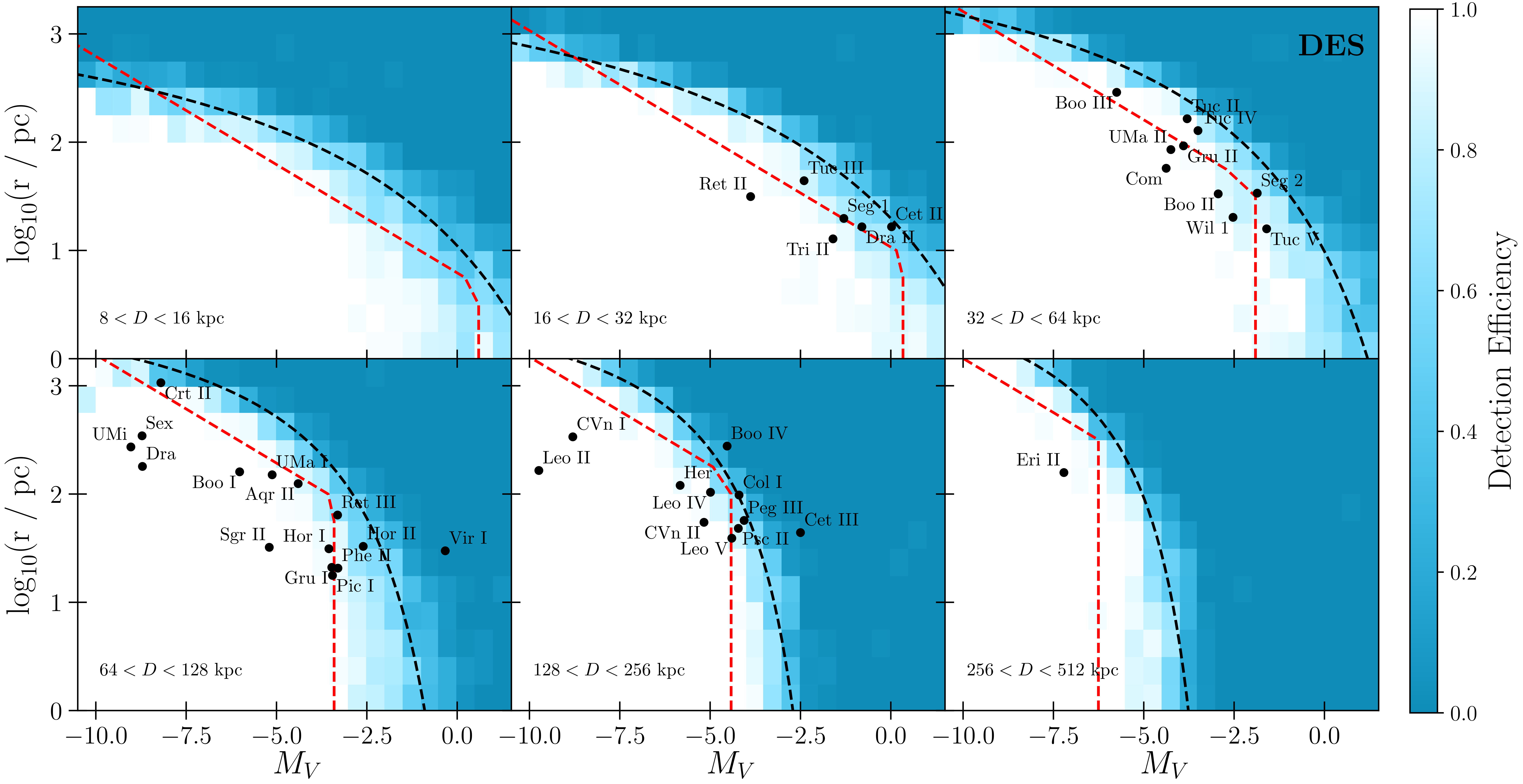
DES & Pan-STARRS Survey Selection Functions



- Rigorous satellite search over ~75% of the sky
- Masks for dusty regions, background galaxies, etc.
- 17/18 (DES), 19/31 (PS1) satellites recovered by two search algorithms
- Selection functions from catalog-level searches for simulated satellites

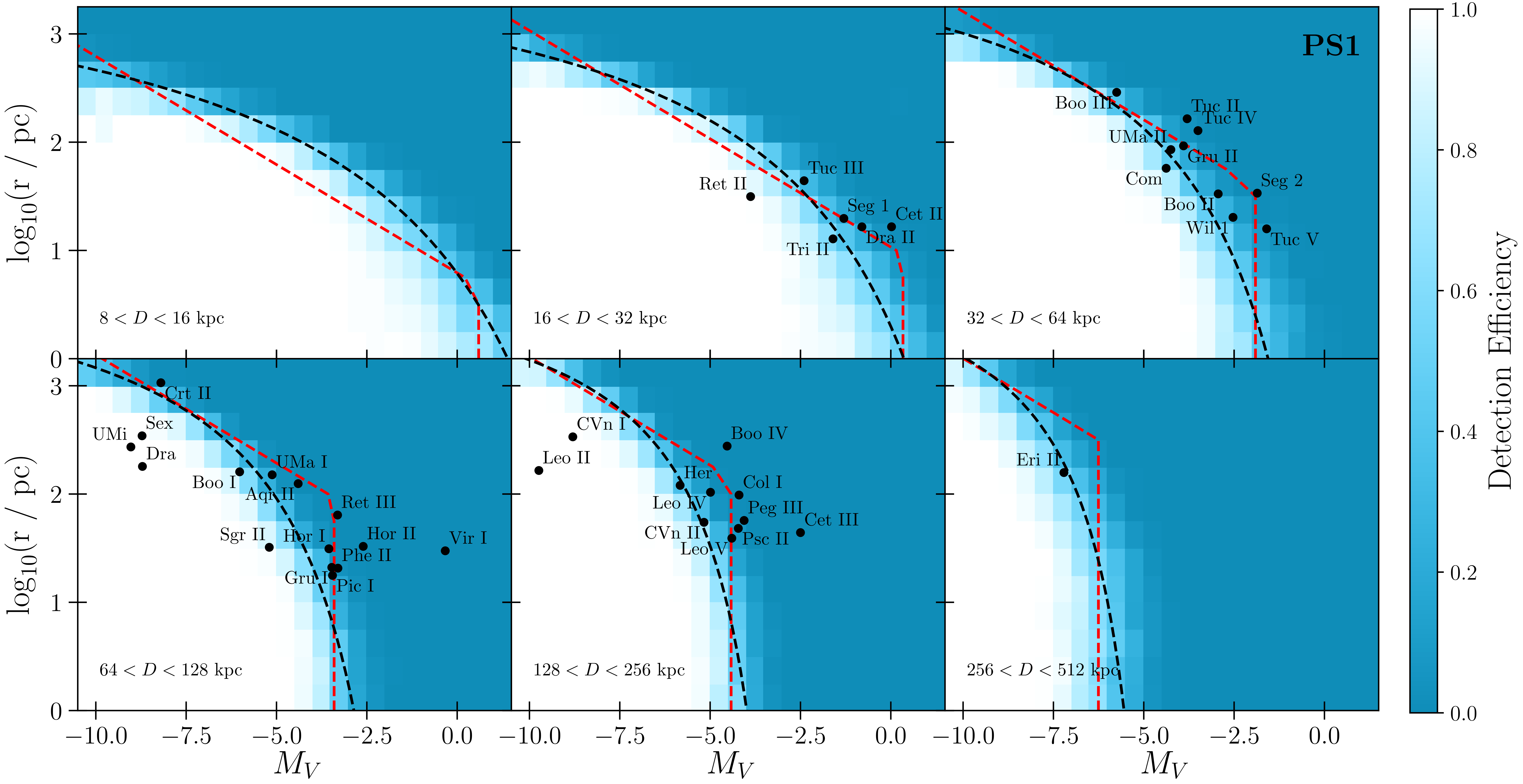
Drlica-Wagner & Bechtol et al. 2020

Observational Selection Functions



Drlica-Wagner & Bechtol et al. 2020

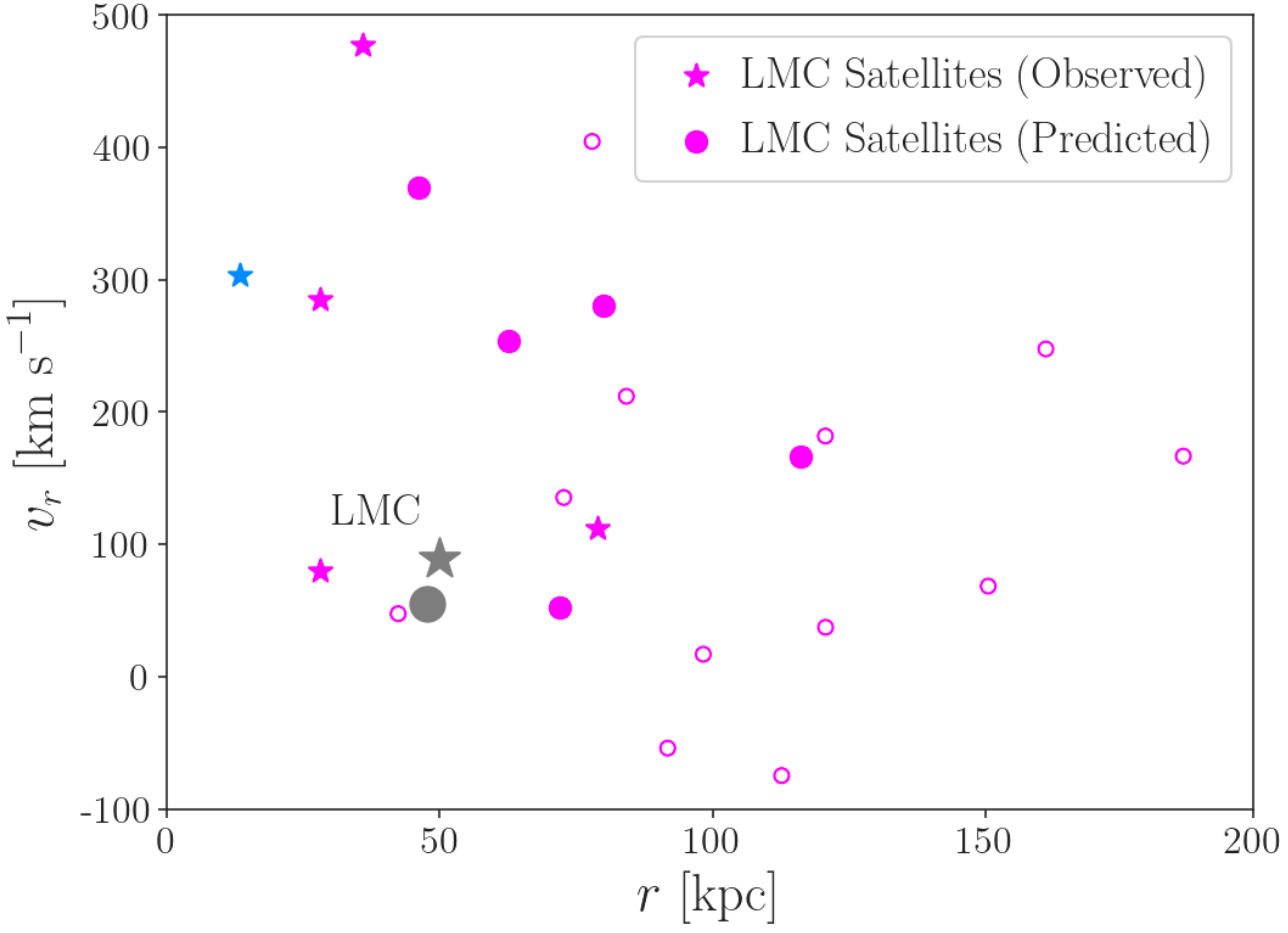
Observational Selection Functions

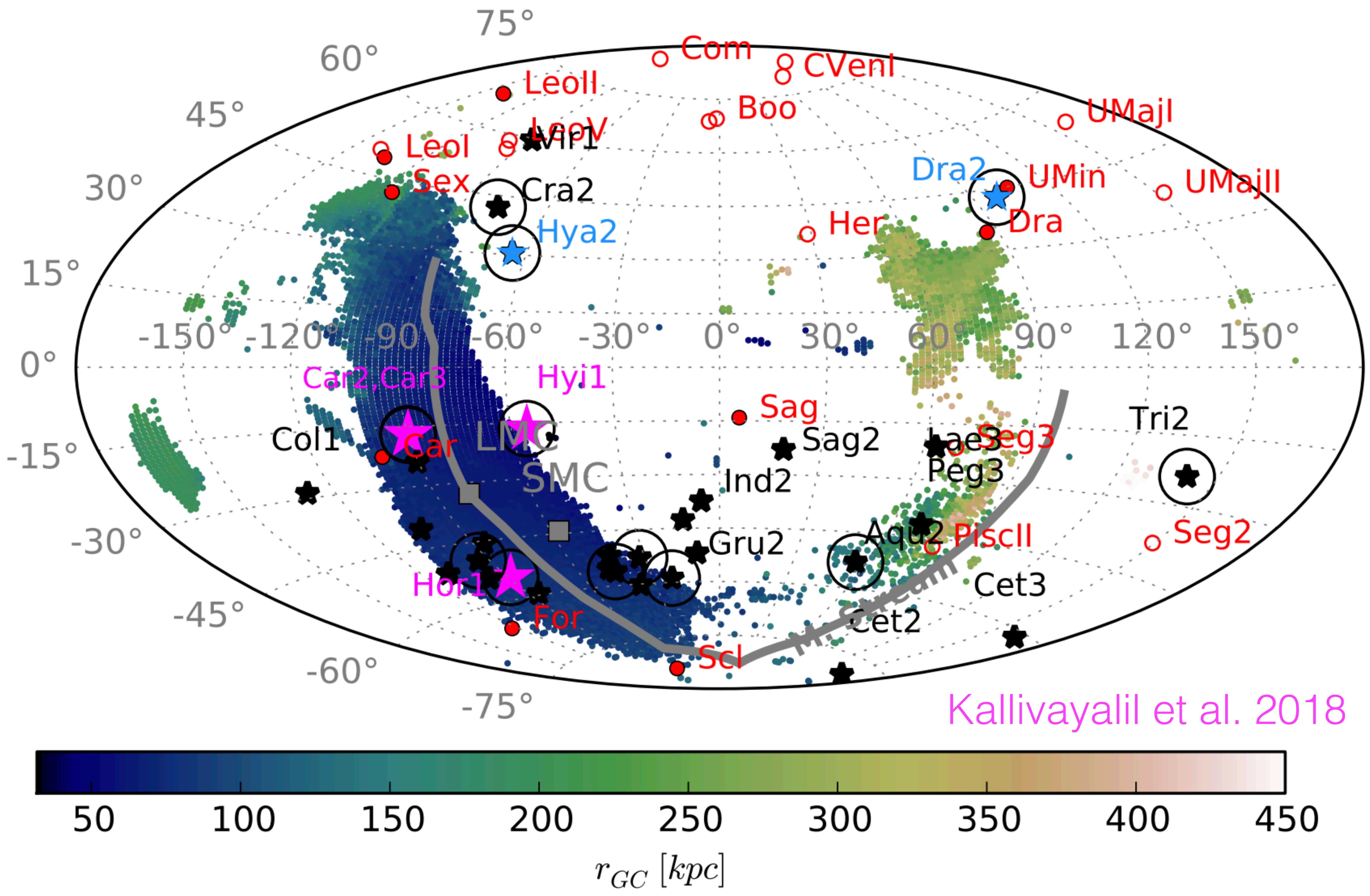


Drlica-Wagner & Bechtol et al. 2020

Satellites of the LMC

Also see Patel et al. 2020!

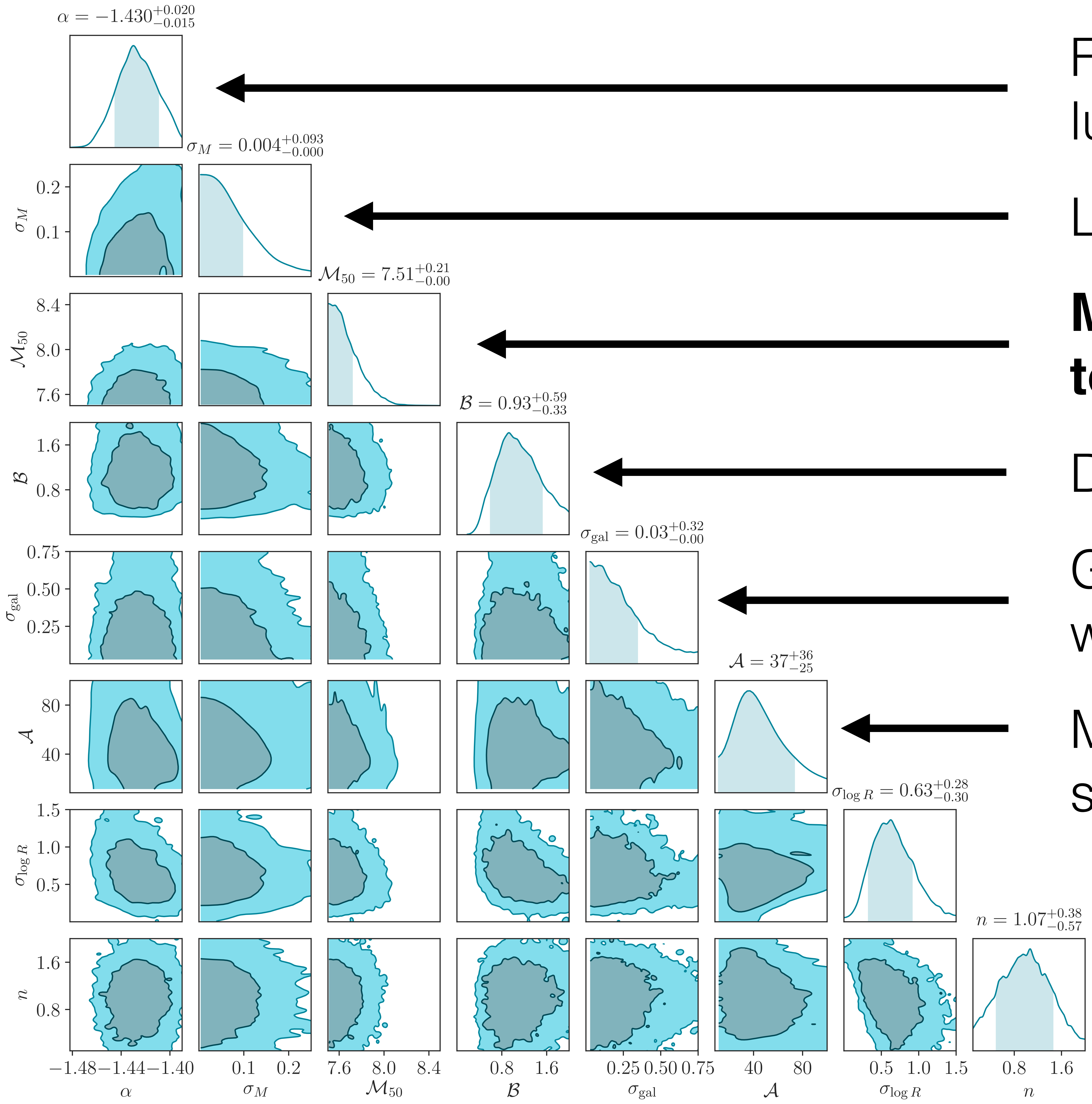




Kallivayalil et al. 2018

Predict 5 (1) currently observed LMC-associated satellites in DES (PS1), consistent with *Gaia* PMs!

EN & Wechsler et al. 2020



Faint-end slope consistent with global luminosity function

Luminosity scatter < 0.2 dex

Minimum halo mass corresponding to observed satellites < 3 x10^{8} M$_{\odot}$

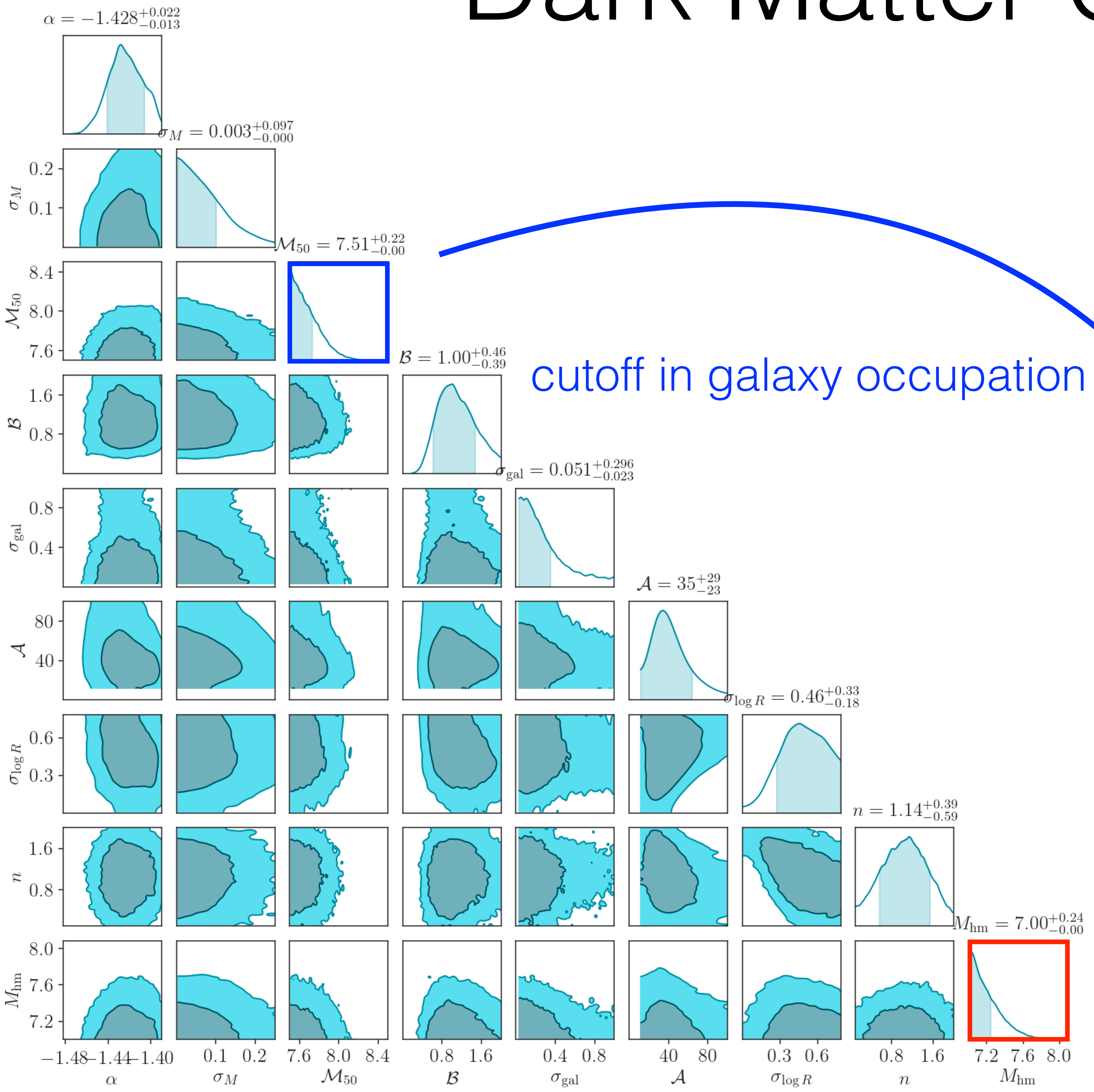
Disruption consistent with FIRE sims

Galaxy occupation fraction consistent with step function

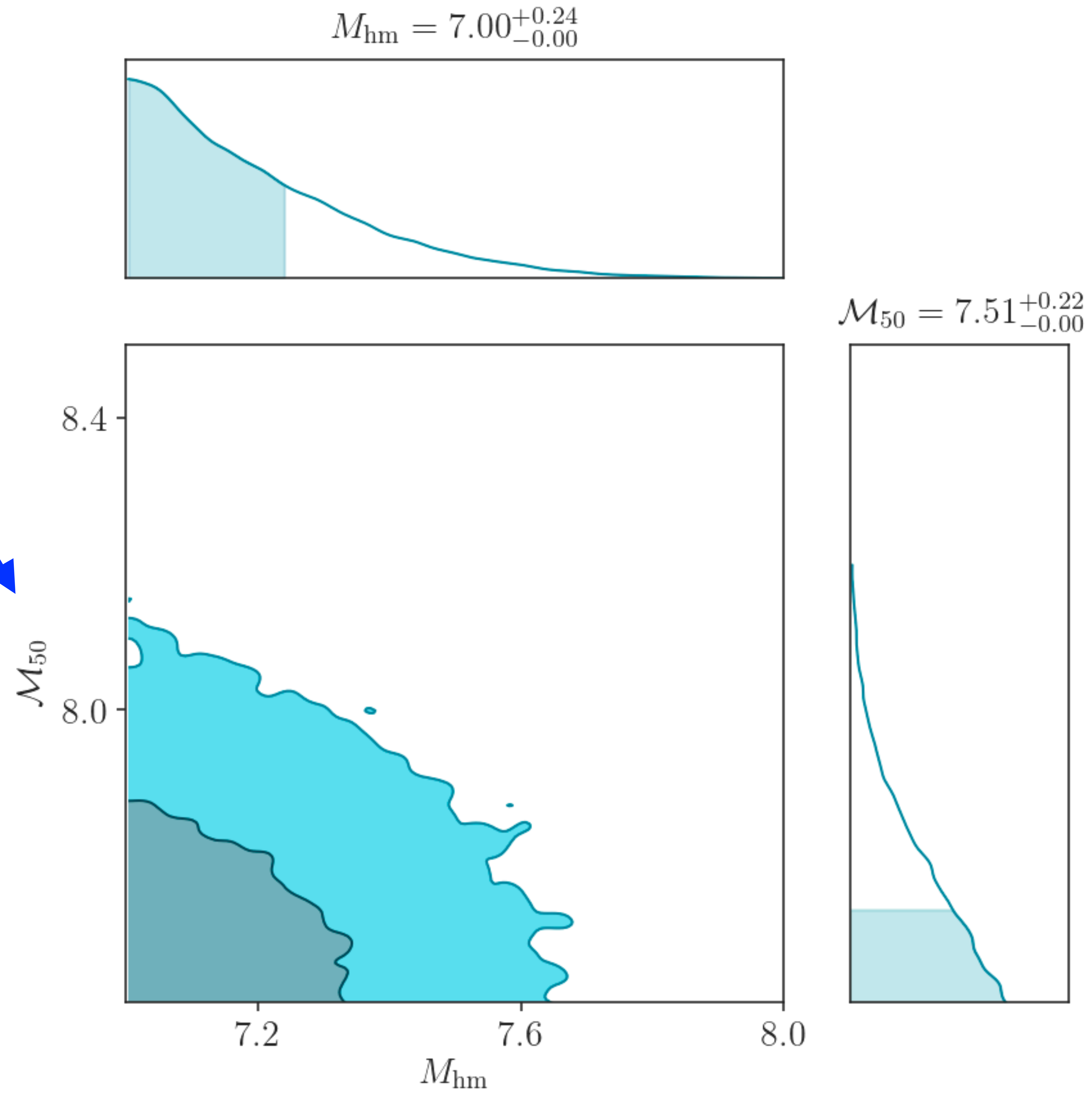
Measurement of amplitude, scatter, slope of galaxy-halo size relation

EN & Wechsler et al. 2020

Dark Matter Constraints



cutoff in galaxy occupation



mass scale of suppression due to dark matter physics

EN & Drlica-Wagner et al. 2020